XENON100 – The new Results

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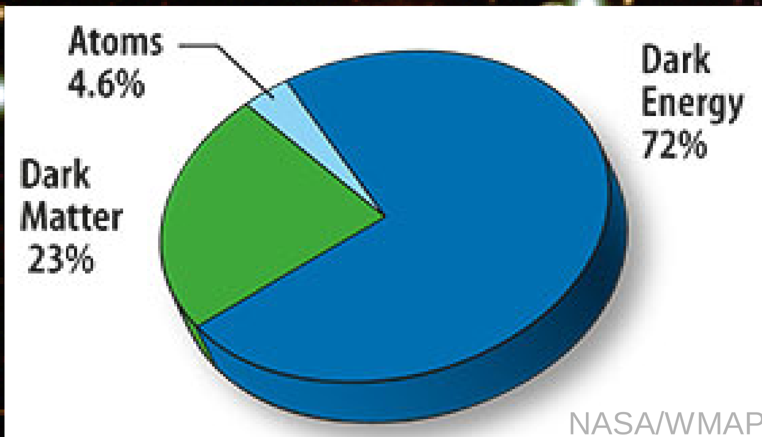
DESY, Hamburg 23.10.2012

www.physik.uzh.ch/groups/groupbaudis/xenon/

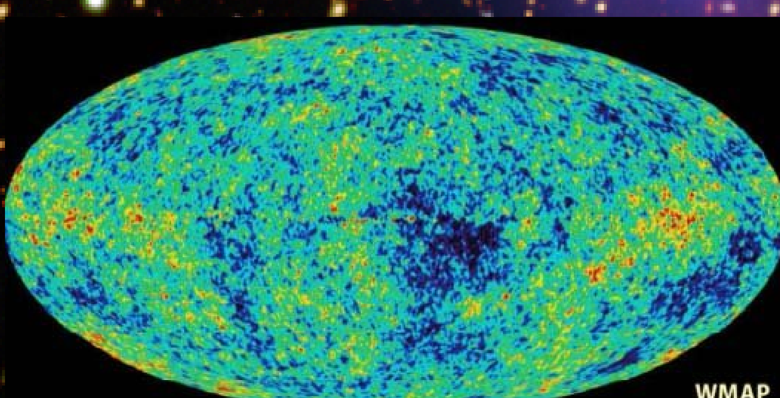
95% of the Universe is dark!

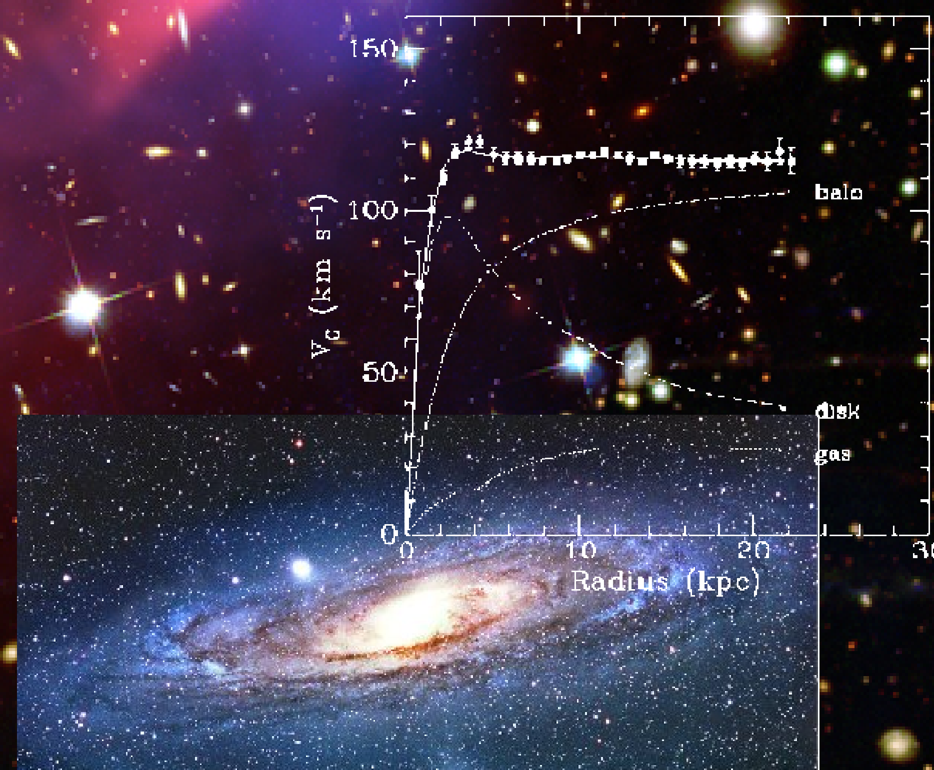
Dark Energy????

Dark Matter: (indirect) **Evidence**

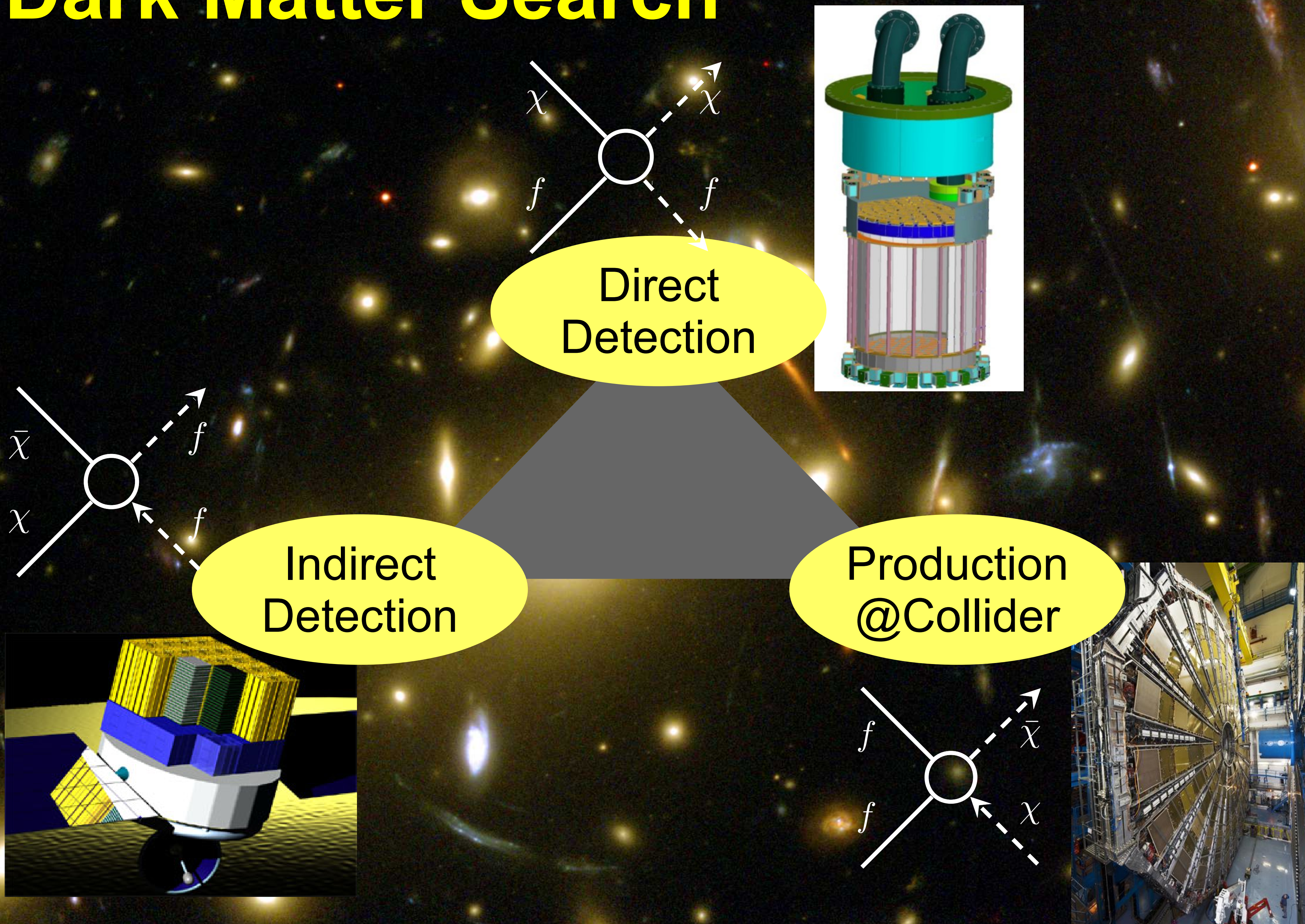


Particle Dark Matter Candidates:
– WIMP → *„WIMP miracle“*
– Axion
– SuperWIMPs
– sterile neutrinos
– WIMPless dark matter
– Gravitino
– ...



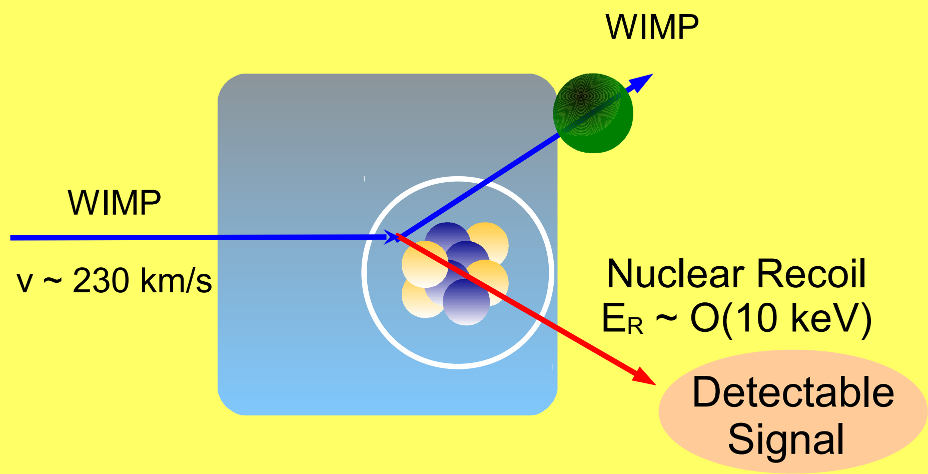


Dark Matter Search



Direct WIMP Search

Elastic Scattering of WIMPs off target nuclei
→ nuclear recoil



Recoil Energy: $$E_r = \frac{|\vec{q}|^2}{2m_N} = \frac{\mu^2 v^2}{m_N}(1 - \cos\theta) \sim \mathcal{O}(10\ \mathrm{keV})$$

Event Rate: $$R \propto N \frac{\rho_\chi}{m_\chi} \langle \sigma_{\chi - N} \rangle$$

N	number of target nuclei
ρ_χ / m_χ	local WIMP density
$<\sigma>$	velocity-averaged scatt. X-section

Detector (N) — Local DM Density (ρ_χ / m_χ) — Physics ($\langle \sigma_{\chi - N} \rangle$)

Direct WIMP Search

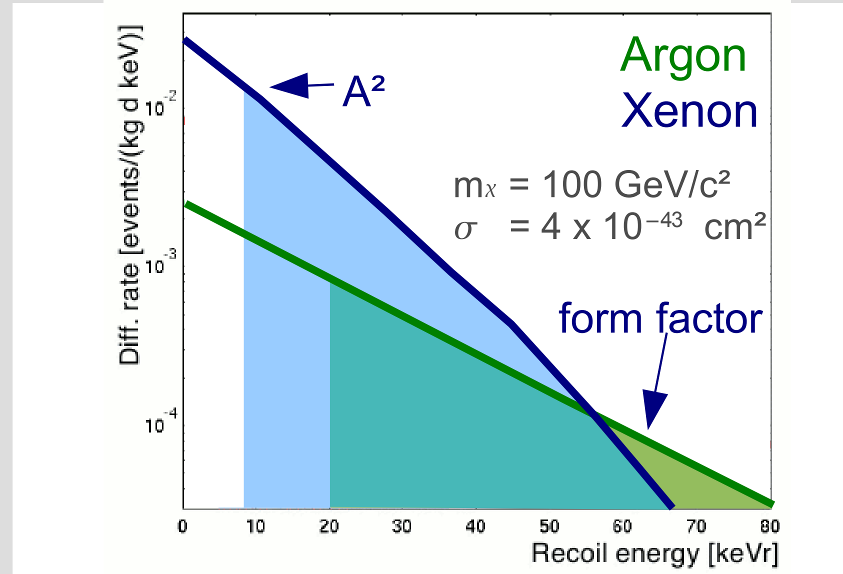
Summary: Tiny Rates

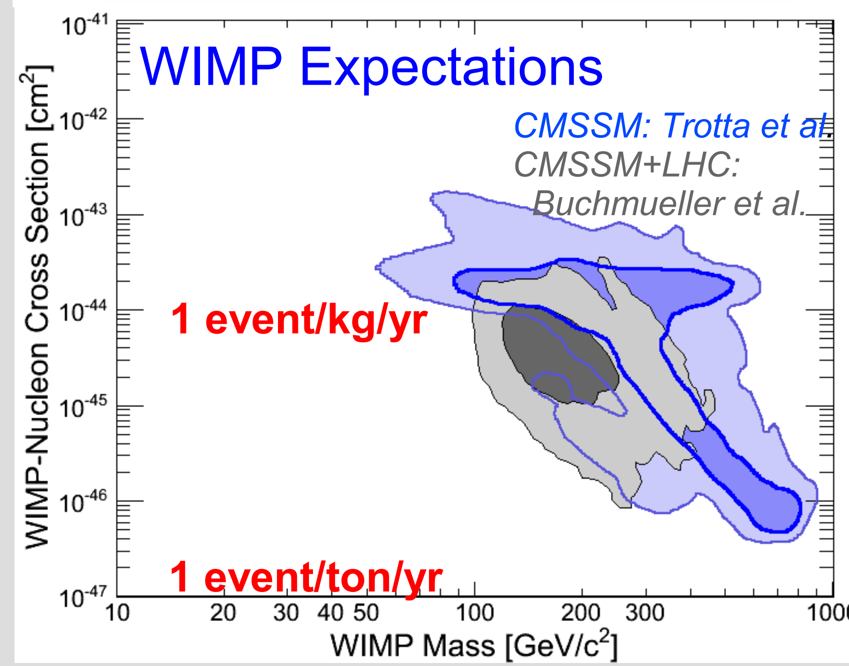
$R < 0.01$ evt/kg/day

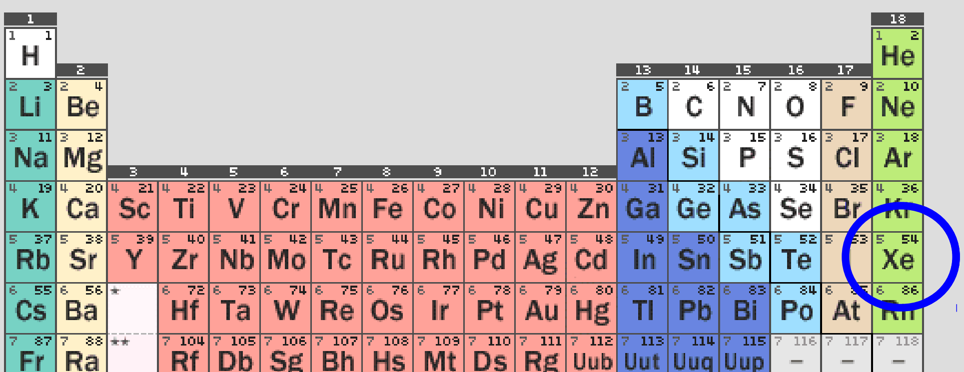
$E_R < 100$ keV

How to build a WIMP detector?

- large total mass, high *A* ✔ for Xe
- low energy threshold ✔ for Xe
- ultra low background ✔ for Xe
- good background discrimination ✔ for Xe







Backgrounds

Experimental Sensitivity

without background: $\propto (mt)^{-1}$
with background: $\propto (mt)^{-1/2}$

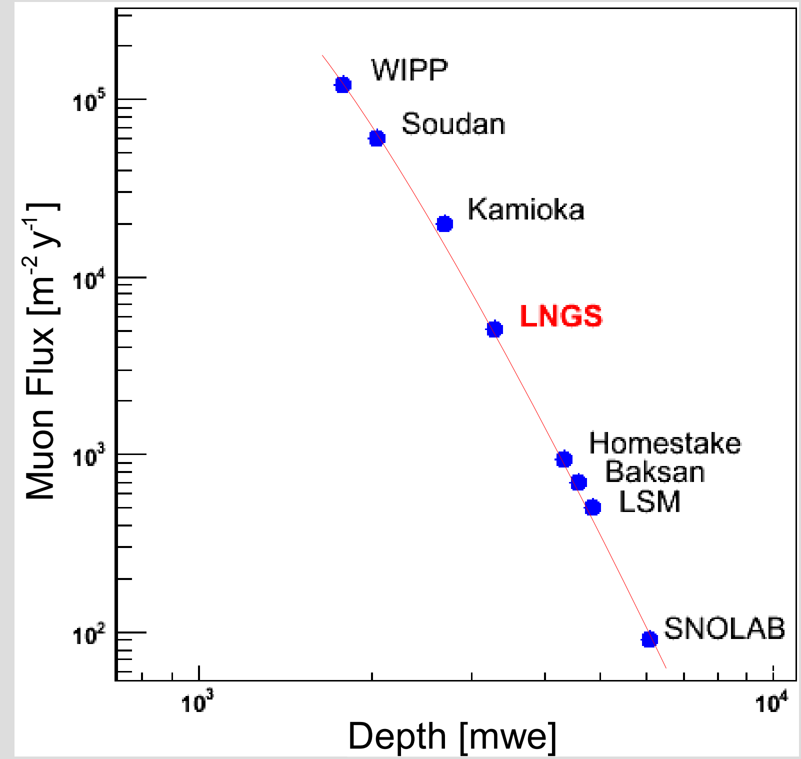
Background Sources

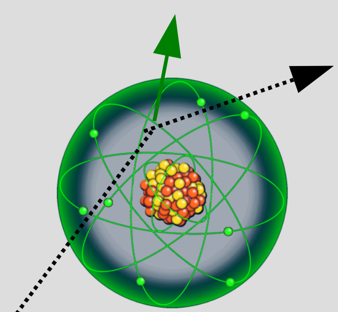
environment: U, Th chains, K

$$^{238}U \xrightarrow[\alpha]{} {}^{234}Th \xrightarrow[\beta]{} {}^{234m}Pa \xrightarrow[\beta]{} {}^{234}U \xrightarrow[\alpha]{} {}^{230}Th \xrightarrow[\alpha]{} {}^{226}Ra \xrightarrow[\alpha]{} {}^{222}Rn \xrightarrow[\alpha]{} {}^{218}Po \ldots$$

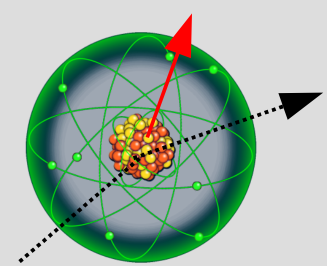
$$^{232}Th \xrightarrow[\alpha]{} {}^{228}Ra \xrightarrow[\beta]{} {}^{228m}Ac \xrightarrow[\beta]{} {}^{228}Th \xrightarrow[\alpha]{} {}^{224}Ra \xrightarrow[\alpha]{} {}^{220}Rn \xrightarrow[\alpha]{} {}^{216}Po \ldots$$

- γ and β decays (electronic recoil)
- alphas no big problem for LXe
  (technology dependent)
- neutrons from (α,n) and sf in rocks and detector parts
- neutrons from cosmic ray muons





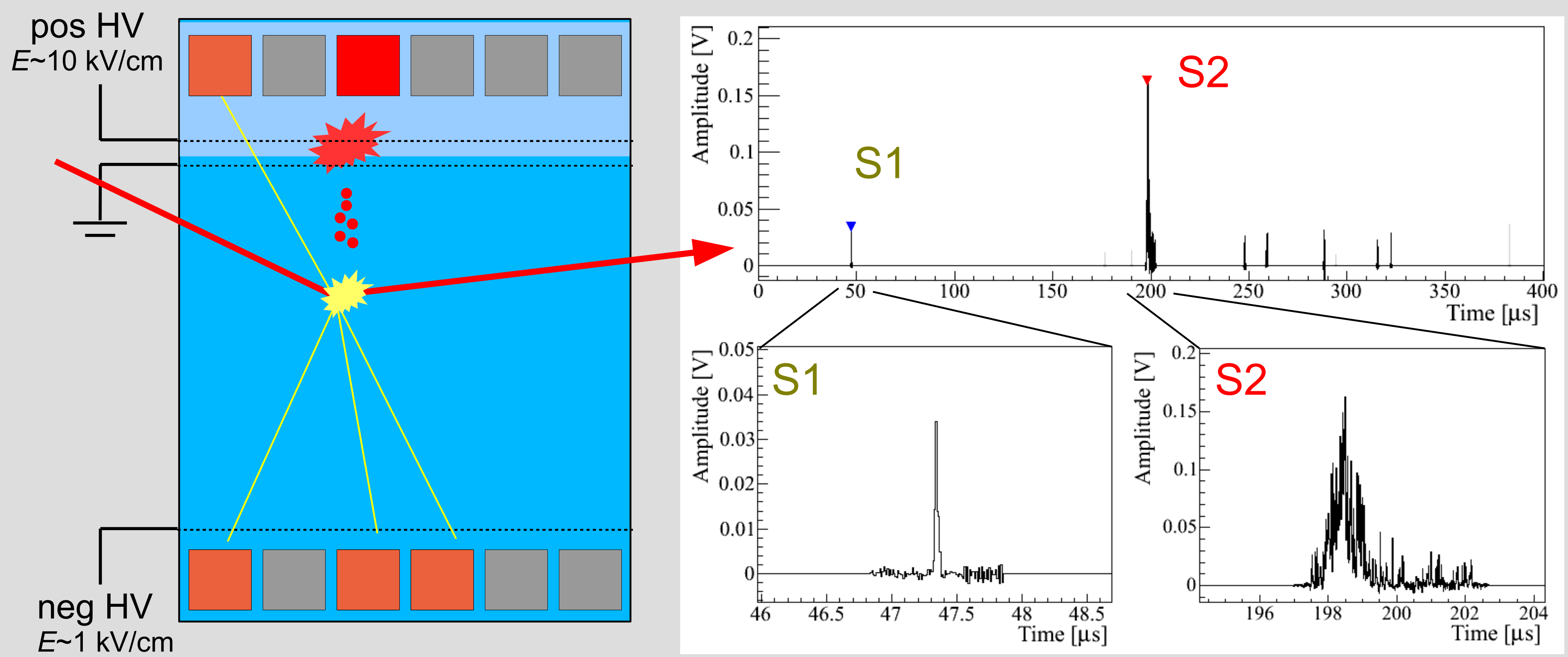
Electronic Recoils
(gamma, beta)



Nuclear Recoils
(neutron, WIMPs)

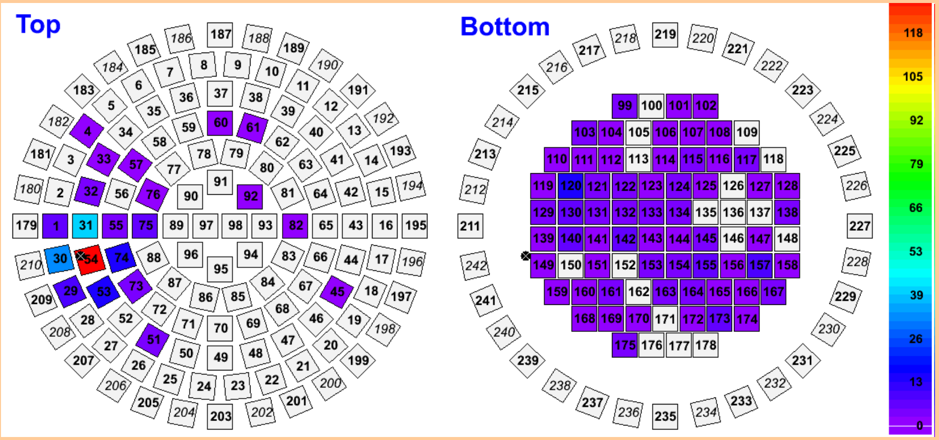
Dual Phase TPC

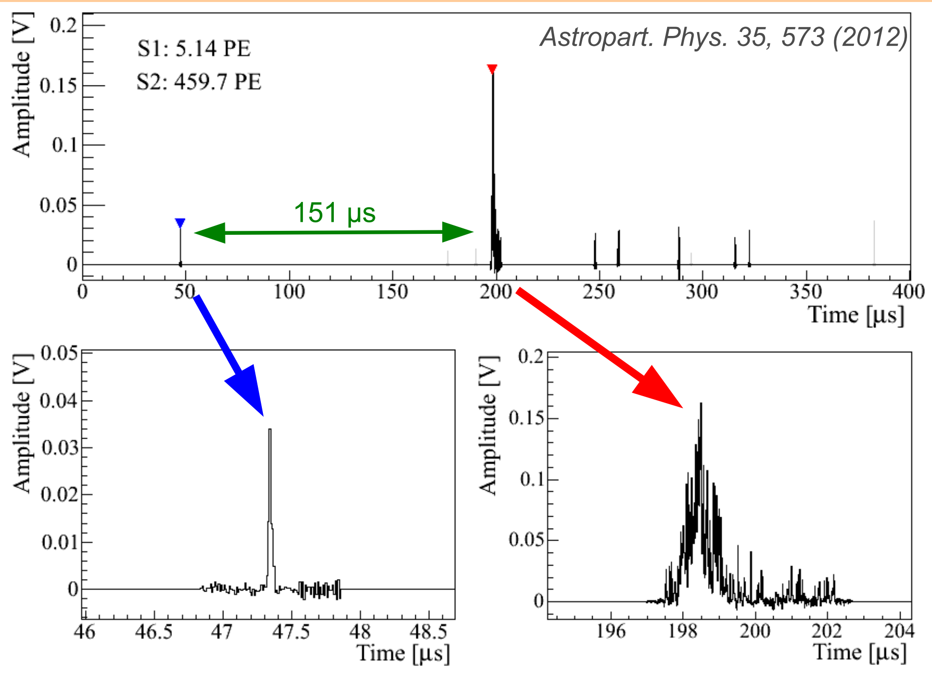
TPC = time projection chamber



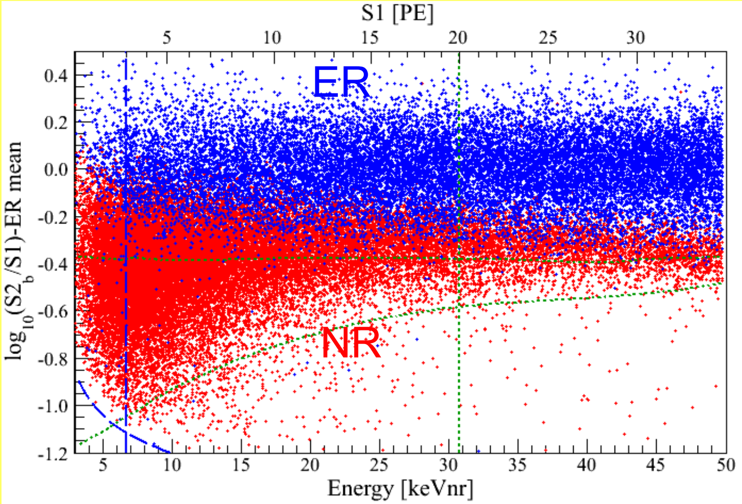
Dual Phase TPC

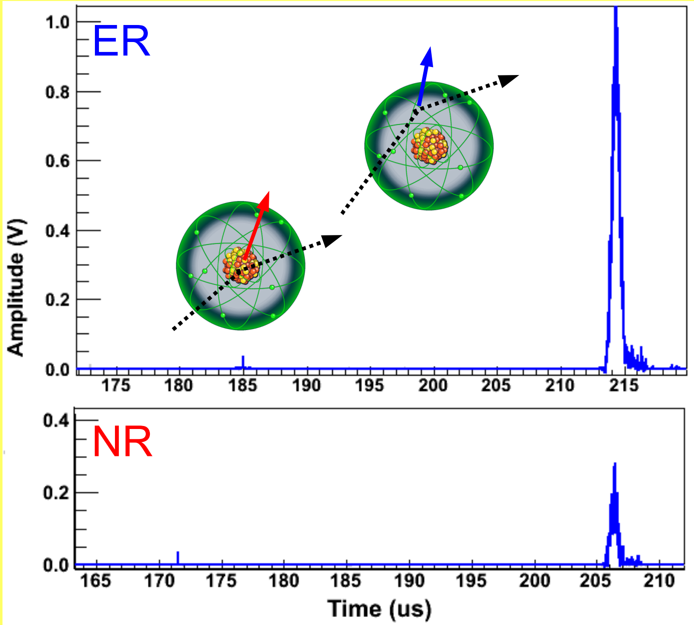
3d Vertex Reconstruction





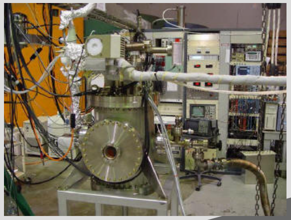
Signal/Background Discrimination





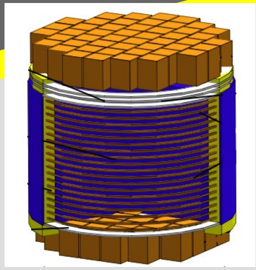
The XENON program

XENON: A phased WIMP search program

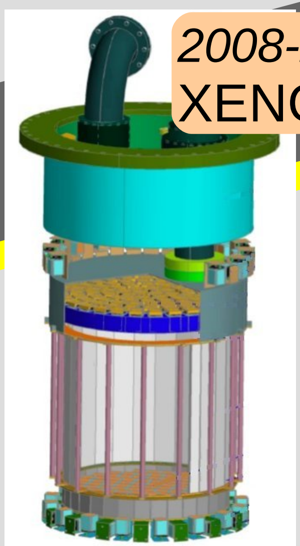


XENON R&D

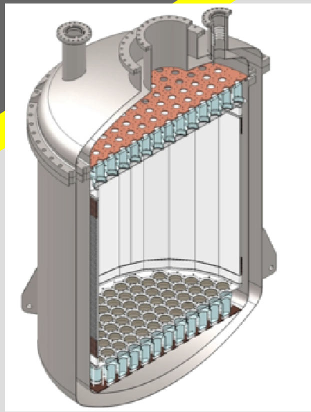
2005-2007: XENON10



2008-2013: XENON100



2010-2017: XENON1T



Columbia *Rice* *UCLA* *U Zürich* *Coimbra* *LNGS* *Mainz* *SJTU*

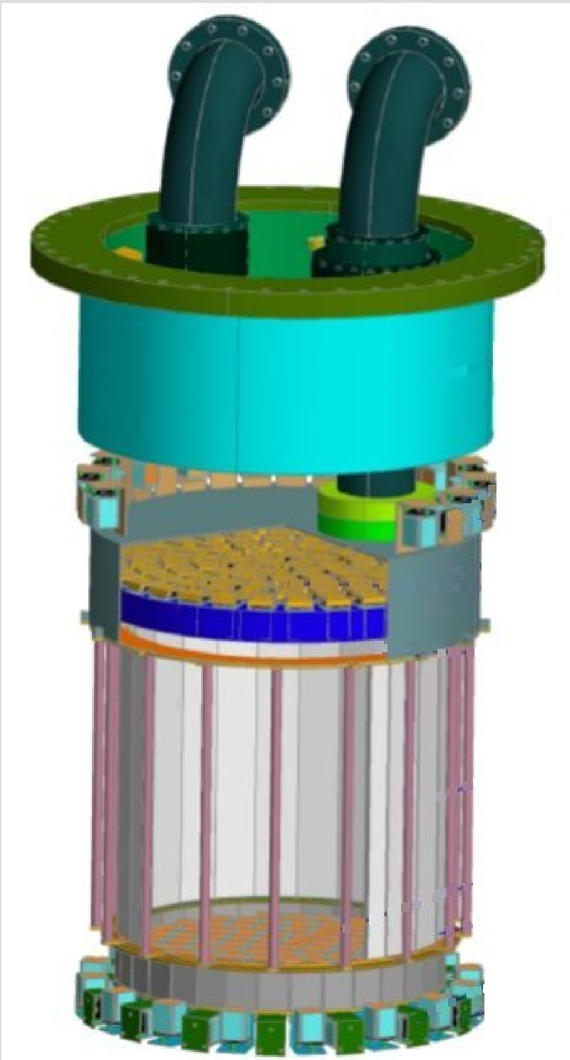
Bologna *MPIK* *NIKHEF* *Purdue* *Subatech* *Münster* *WIS*

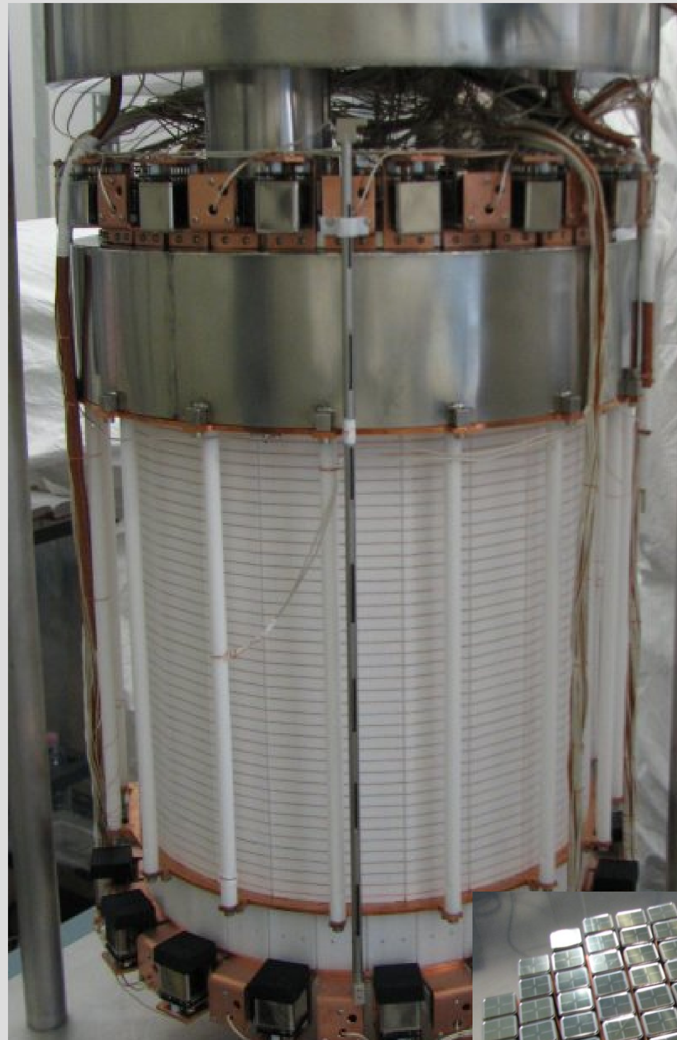
XENON Collaboration

XENON Collaboration Meeting, LNGS, October 2012

XENON100

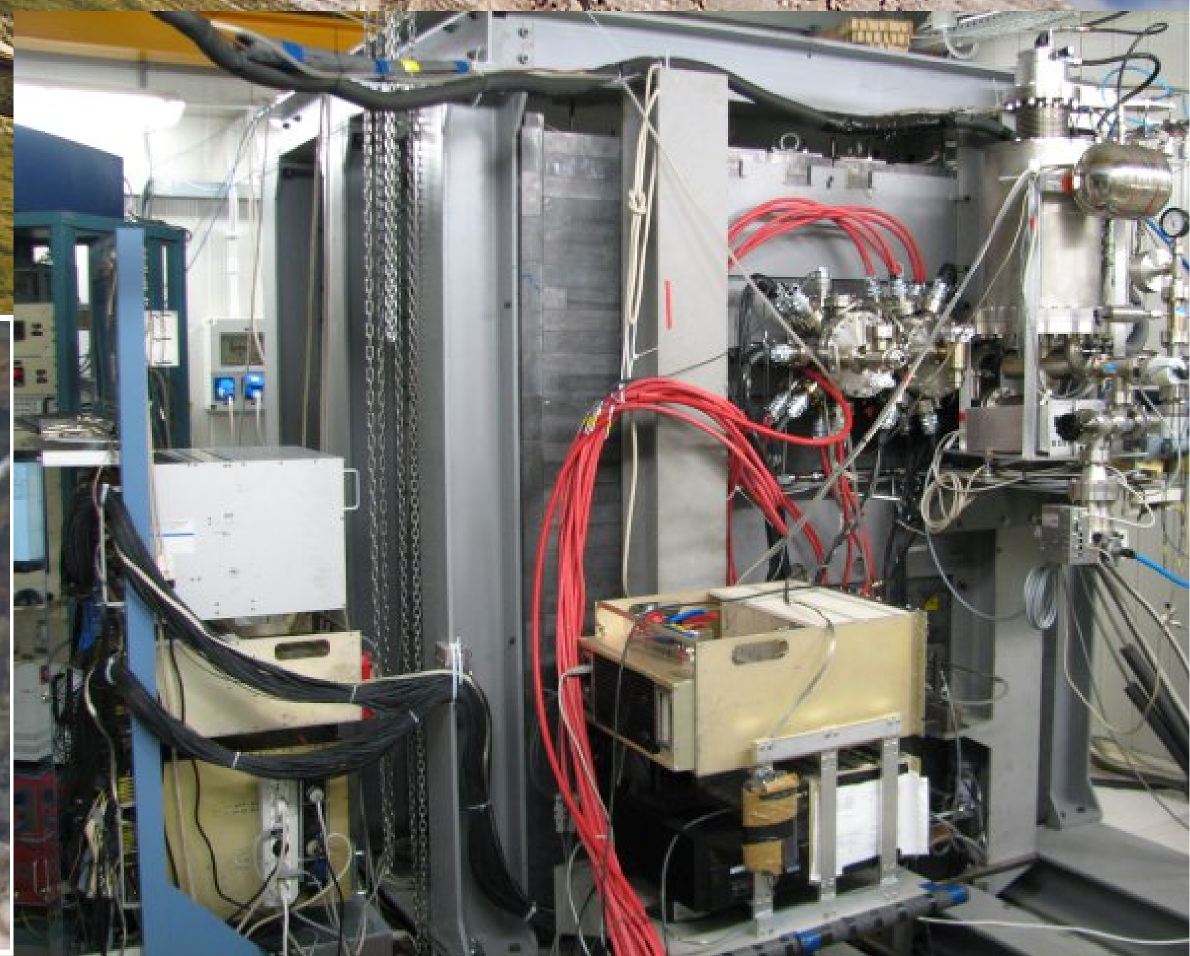
Astropart. Phys. 35, 573 (2012)



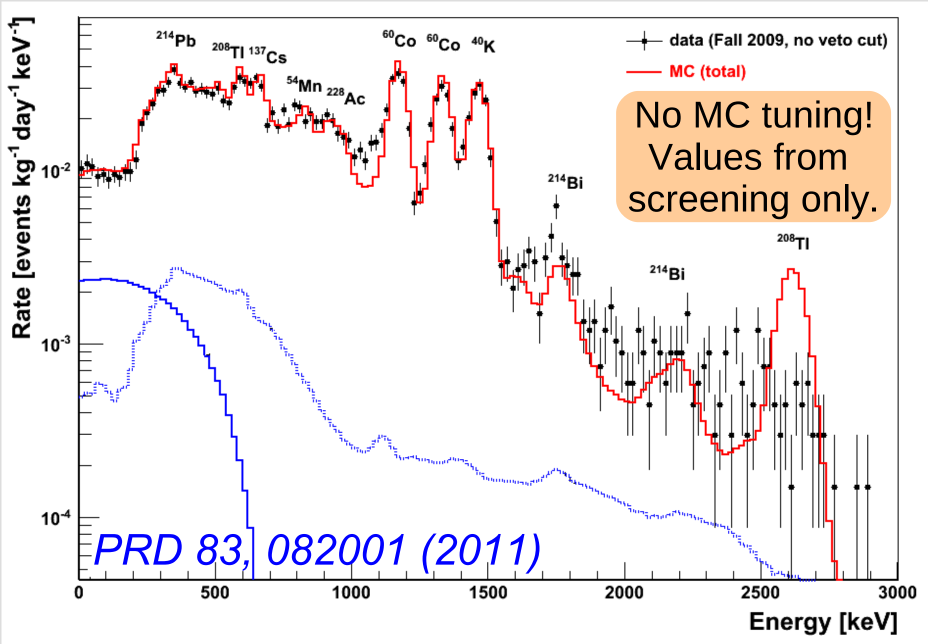


161 kg LXe, 62 kg in target
242 1" x1" PMTs

Laboratori Nazionali del Gran Sasso



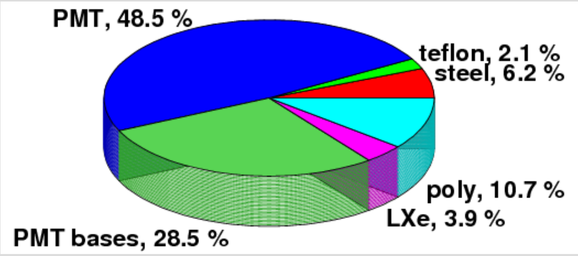
XENON100 Background



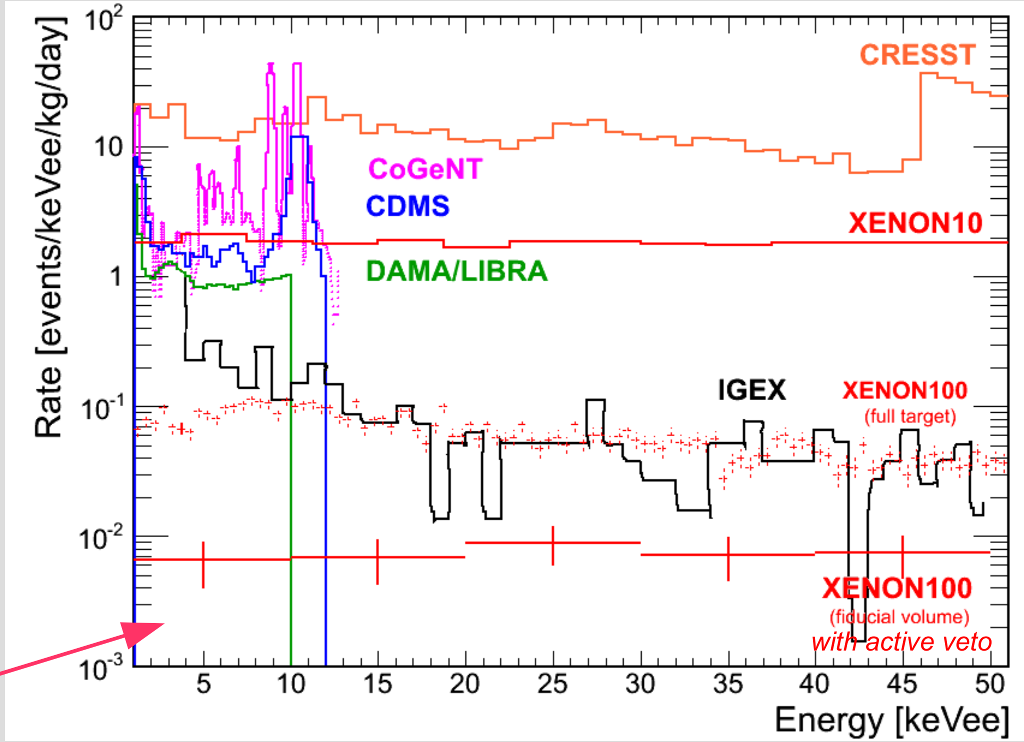
Measured Background in good agreement with MC prediction.

At low energies: Lowest background ever achieved in a Dark Matter Experiment!

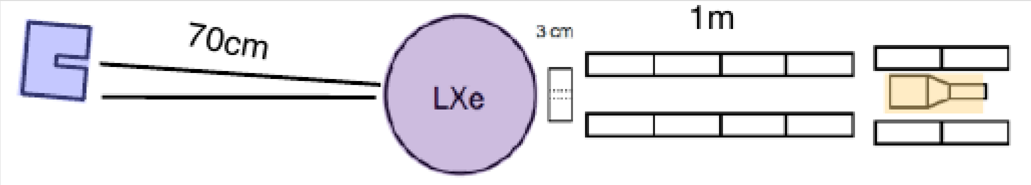
- 30 kg fiducial mass
- active LXe veto not used for this plot
- exploit anti-correlation between light and charge for better ER-energy scale

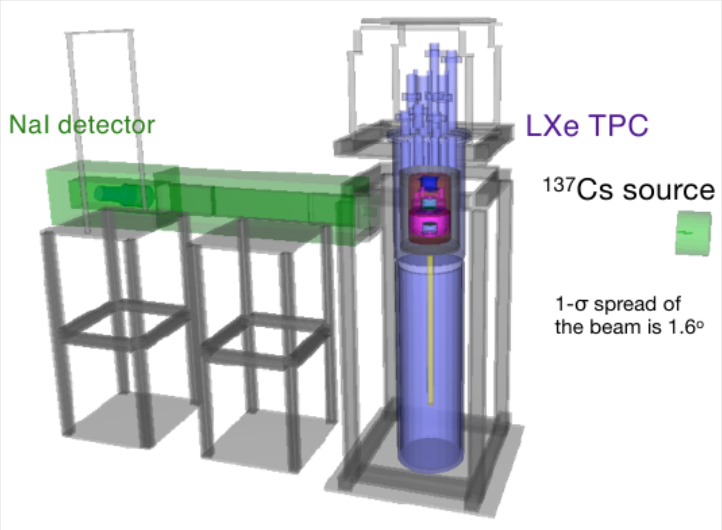


Xenon keVee-Scale not precisely known below 9 keVee

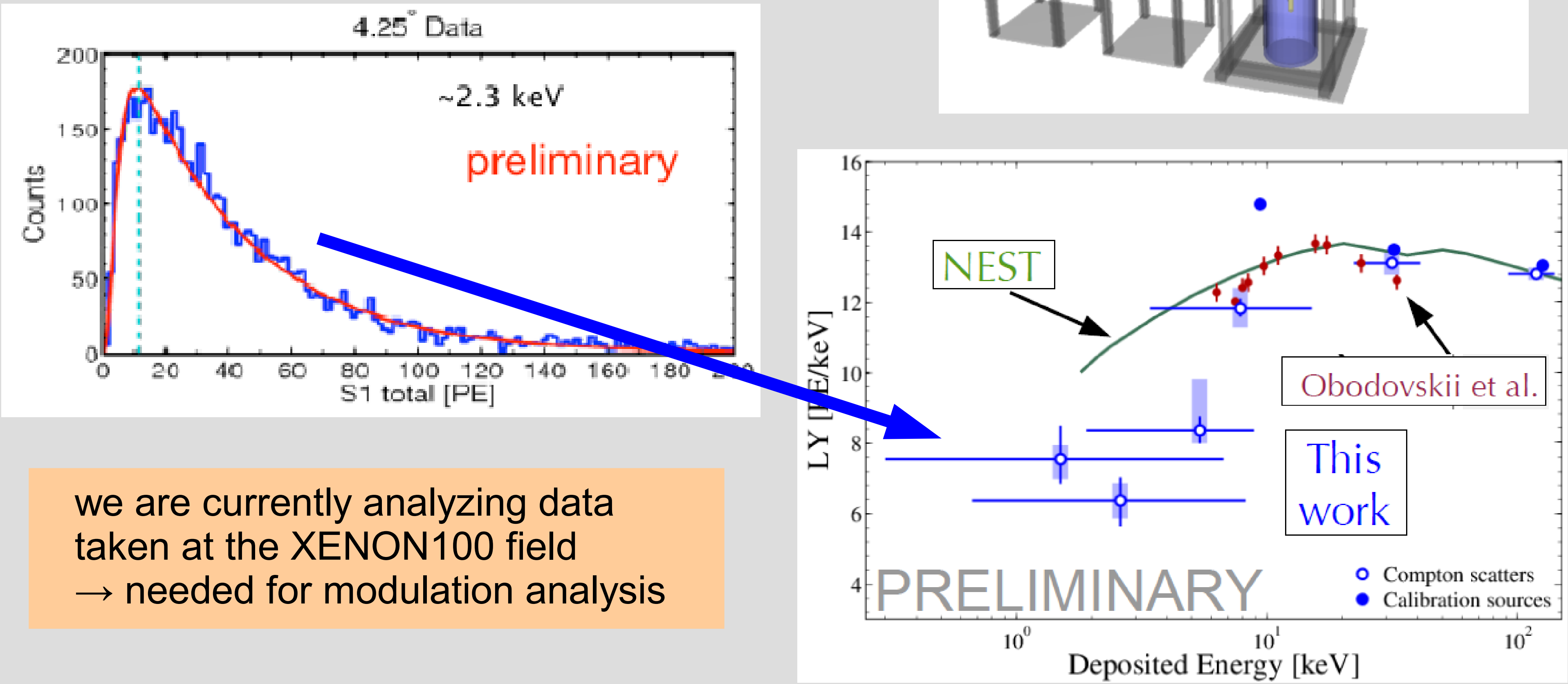


Low Energy Response to ER



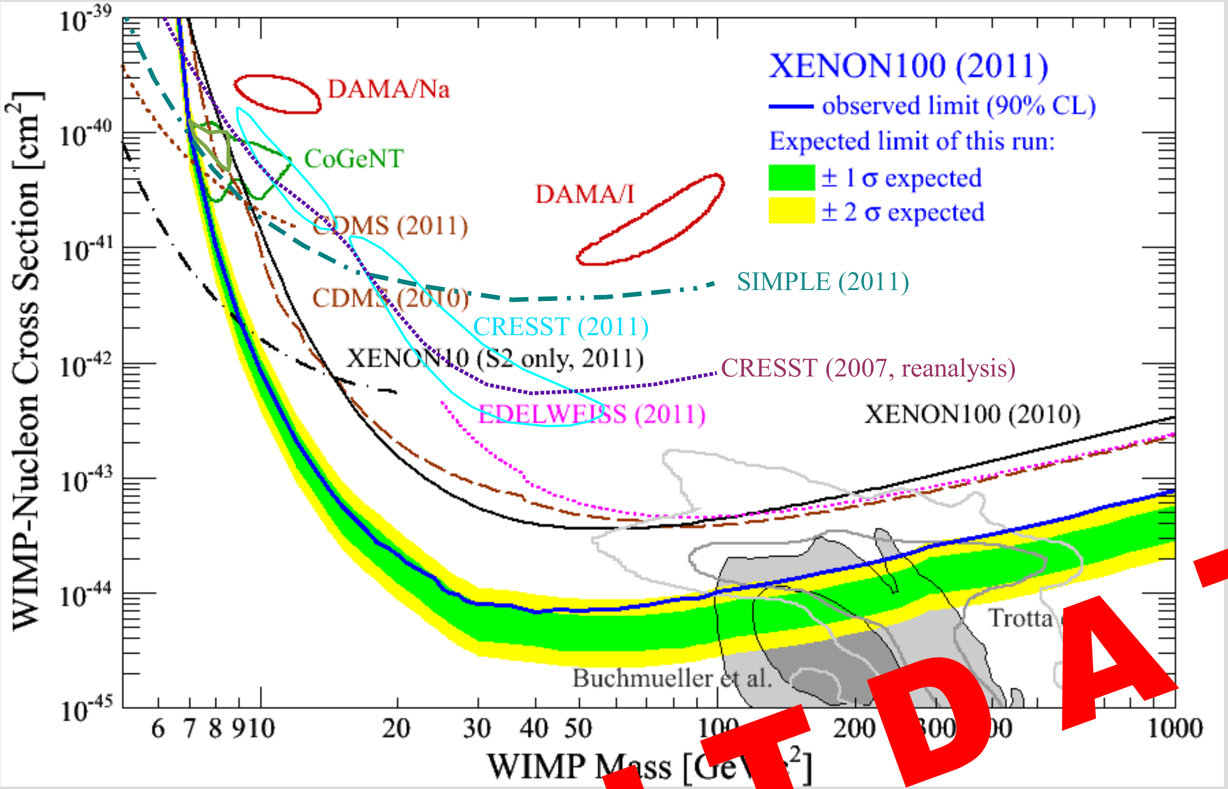


Compton Scatter Measurement at UZH indicates that LXe „sees“ electronic recoil interactions around ~2.3 keV (also with field!!!)



we are currently analyzing data taken at the XENON100 field
→ needed for modulation analysis

(spin-independent) WIMP Limit 2011

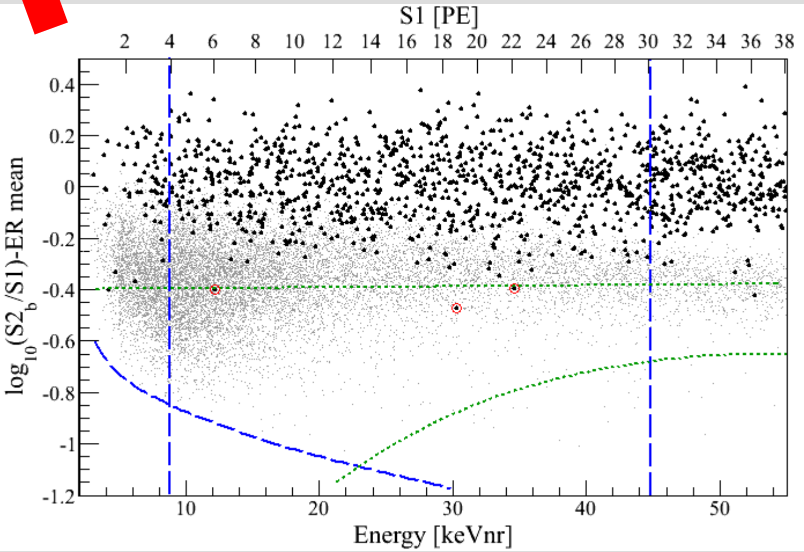


PRL 107, 131302 (2011)

already cited 408x

Limit derived with Profile Likelihood method
PRD 84, 052003 (2011)

OUTDATED!!!



XENON100 sets the most sensitive limit over a large WIMP mass range

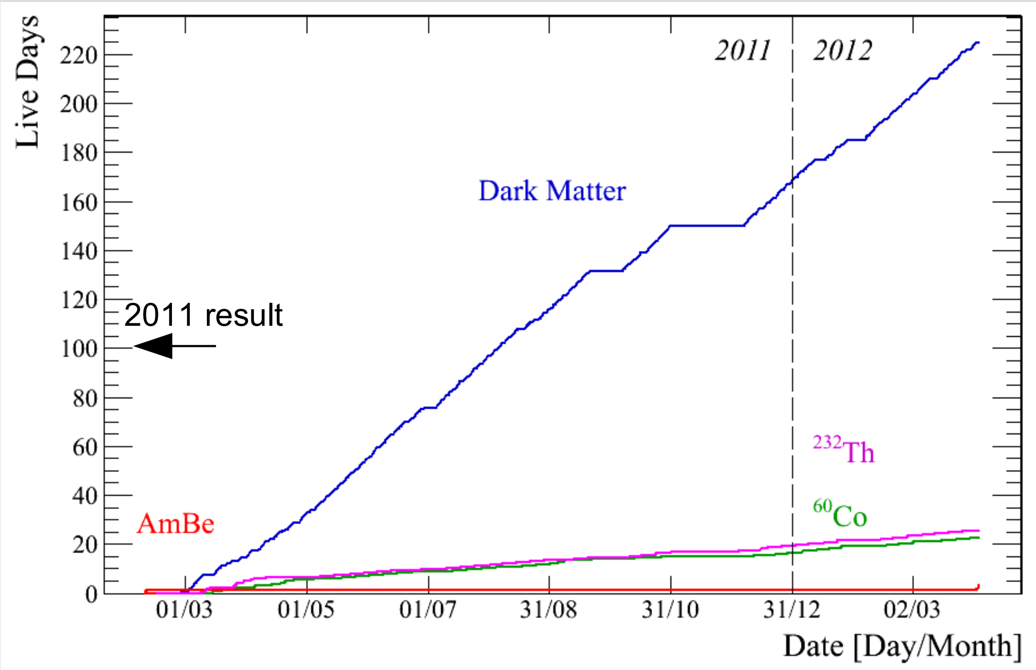
Challenges the CoGeNT, DAMA, CRESST-II signals as being due to light mass WIMPs

XENON100 –
New results of 2012

arXiv:1207.5988, accepted by PRL

Data Taking

Data Collection

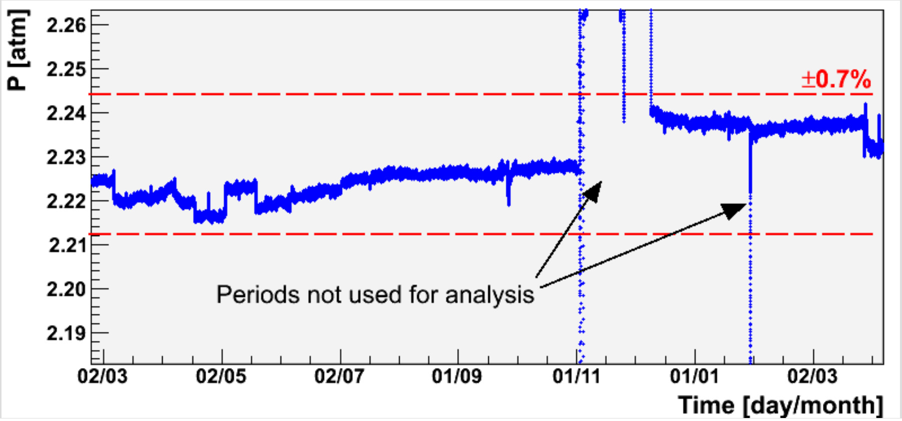


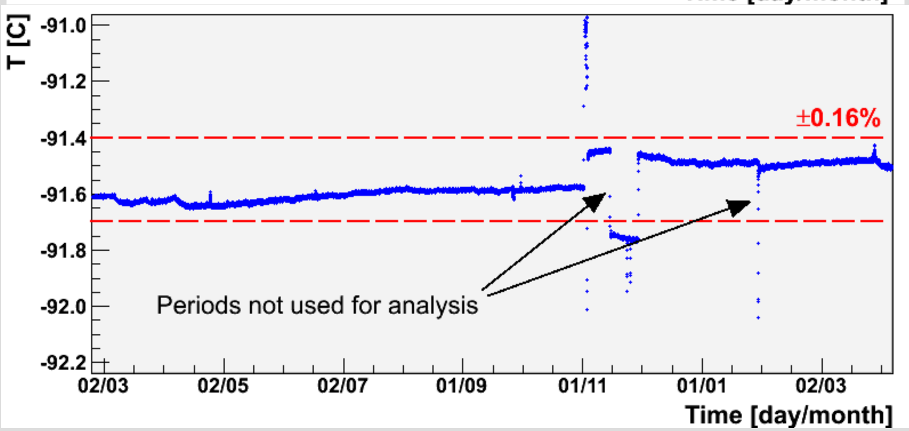
Data taking over 13 months
from Feb 28, 2011 to March 31, 2012
→ full annual cycle

3 interruptions for maintenance

224.56 live days of dark matter data

Stability

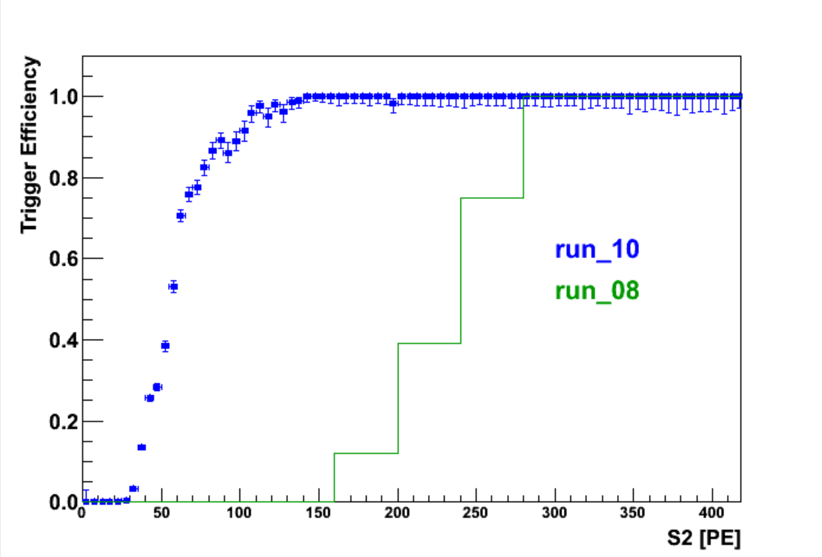


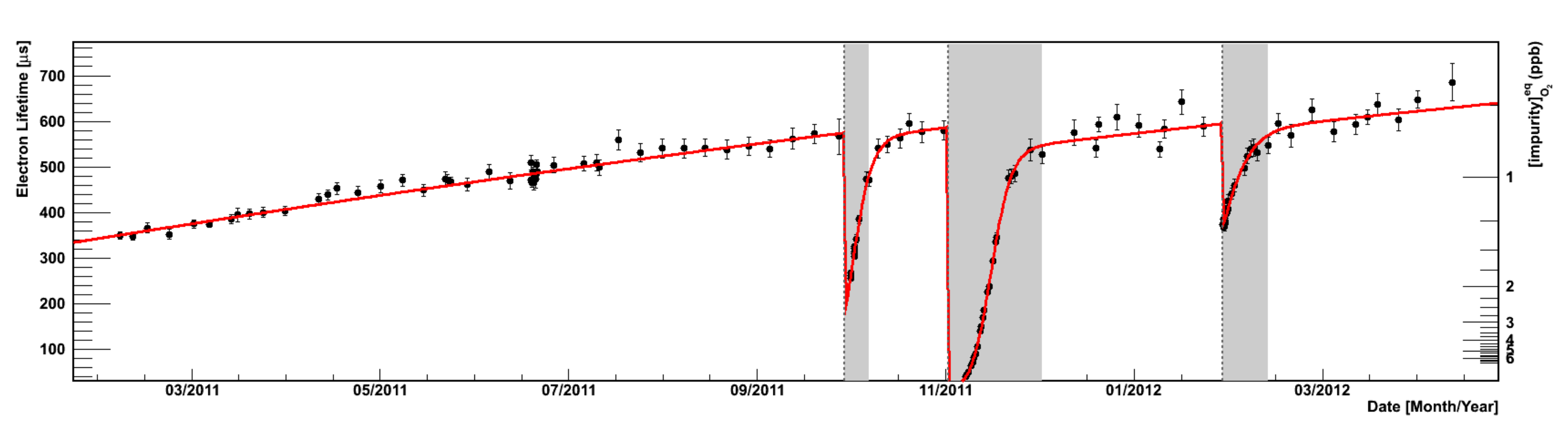


To our knowledge, no large LXe detector has ever been operated under such stable conditions for that long

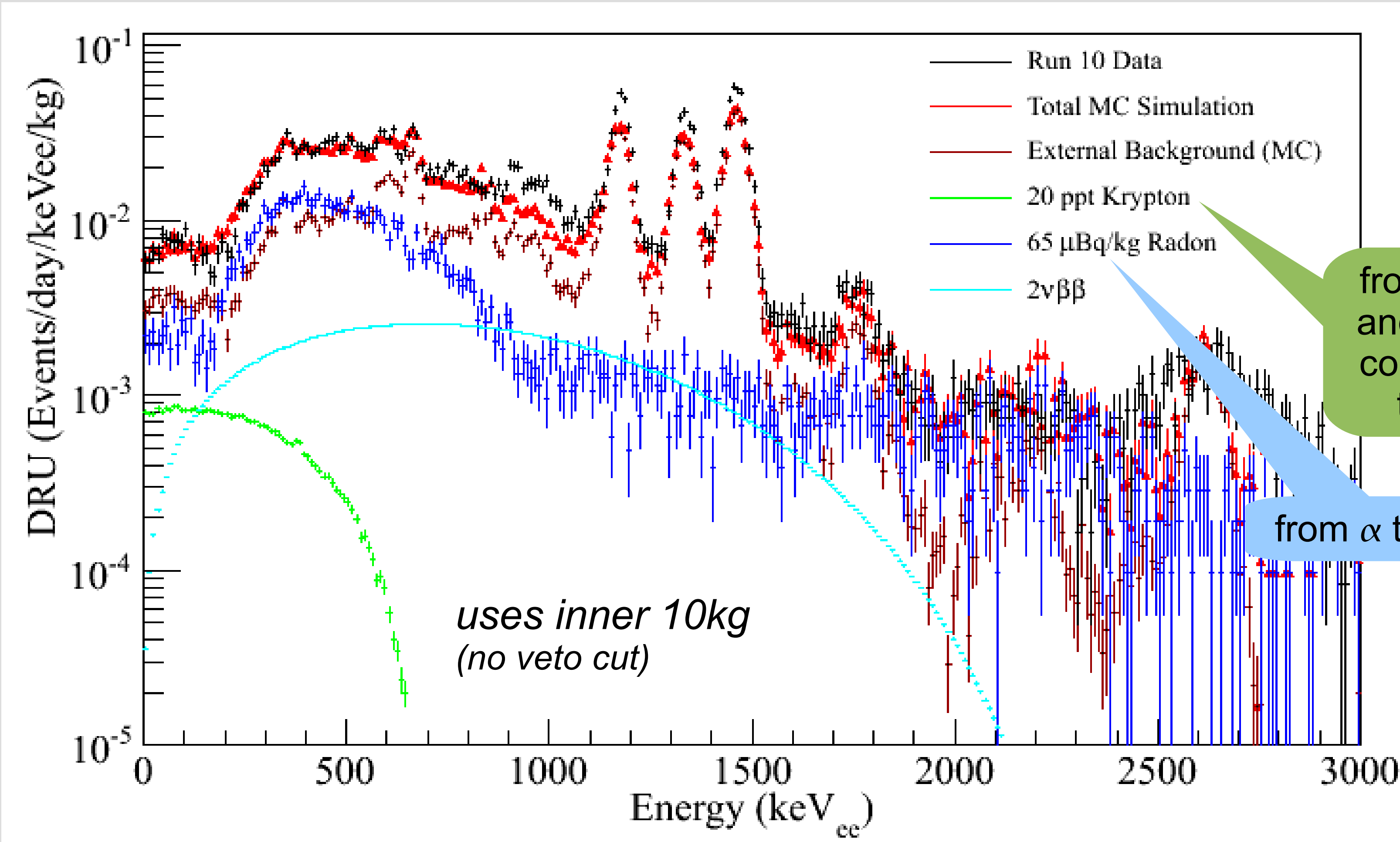
Improvements

- Exposure more than doubled
- Lower threshold
  S2>150 PE, S1>3 PE (6.6 keVr)
- Lower Background
- Much more calibration data
  35x more ER calibration in ROI
  AmBe before and after run
- Higher LXe purity → smaller corrections





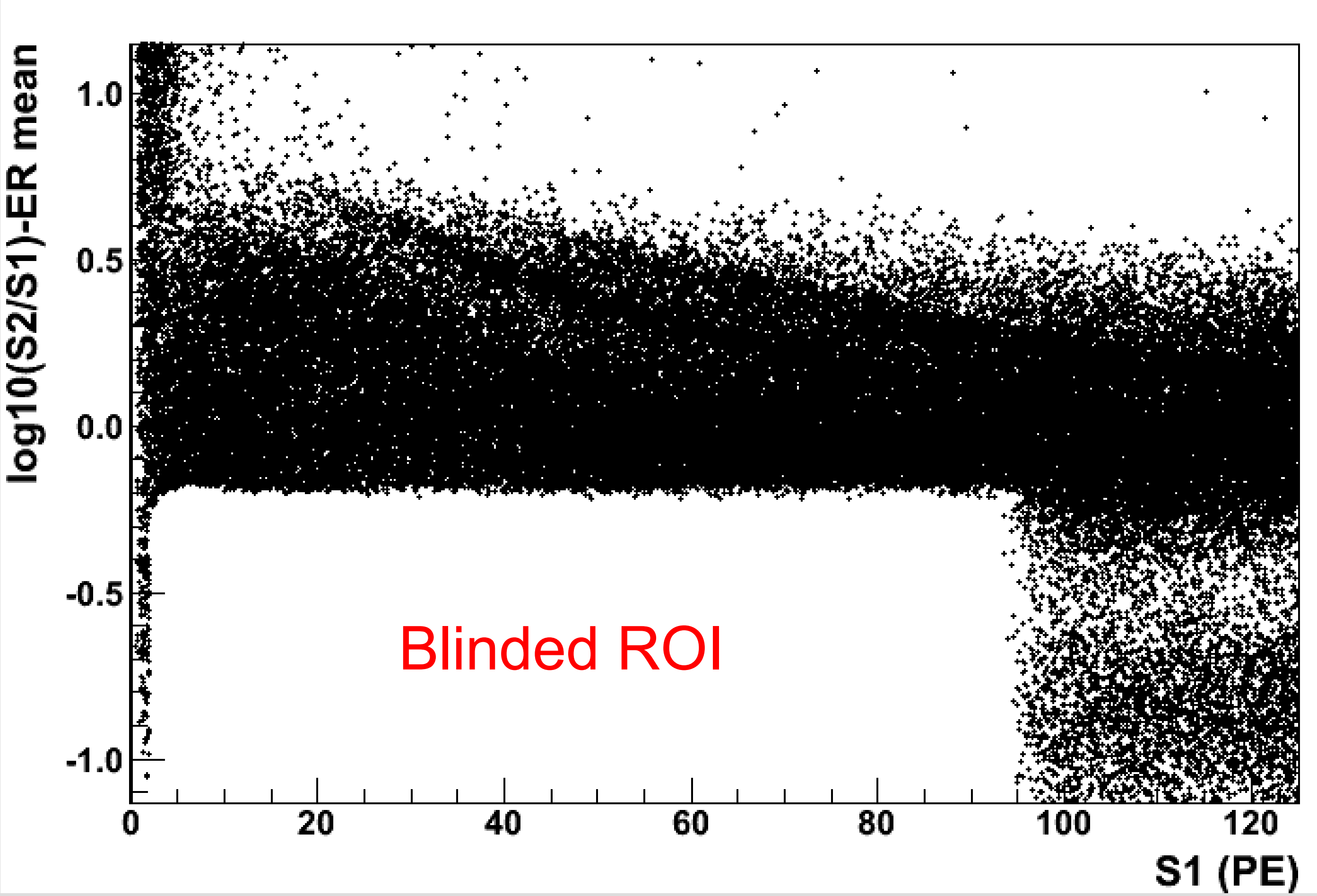
Total ER Background



ER Background: $(5.3 \pm 0.6) \times 10^{-3}$ events/keV/kg/day

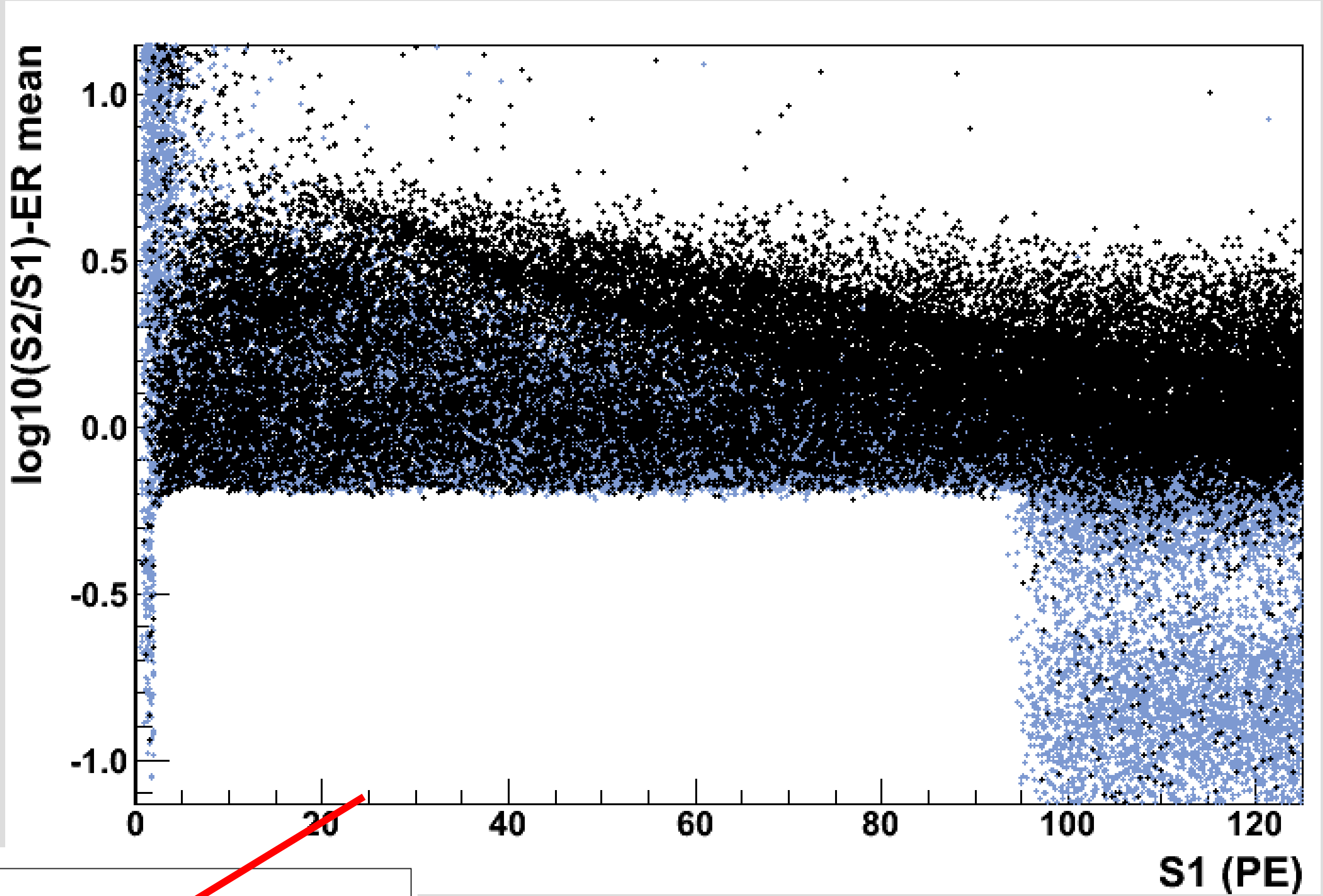
(with active veto) → before discrimination

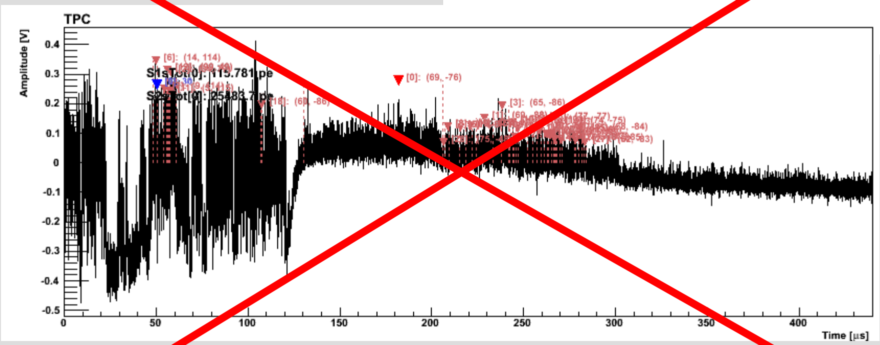
Data Analysis: All data



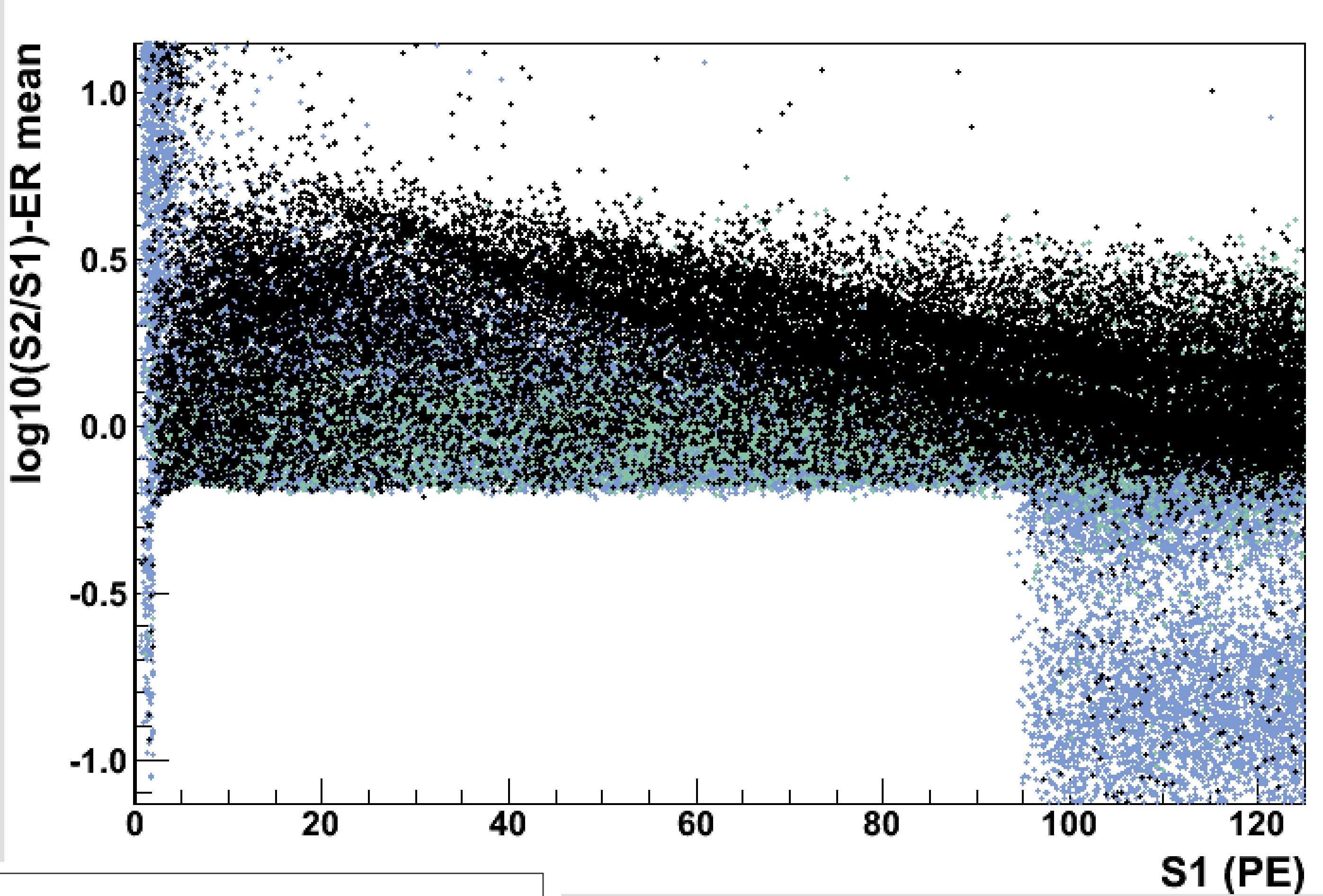
More information on XENON100 data analysis in arXiv:1207.3458

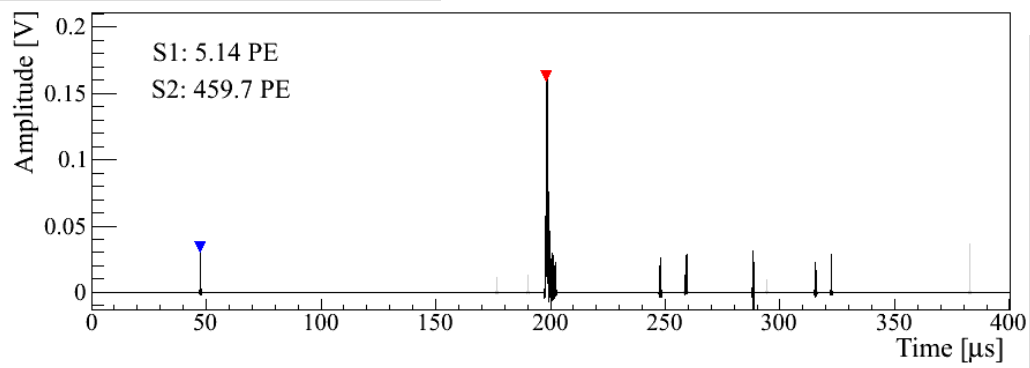
Basic Quality Cuts



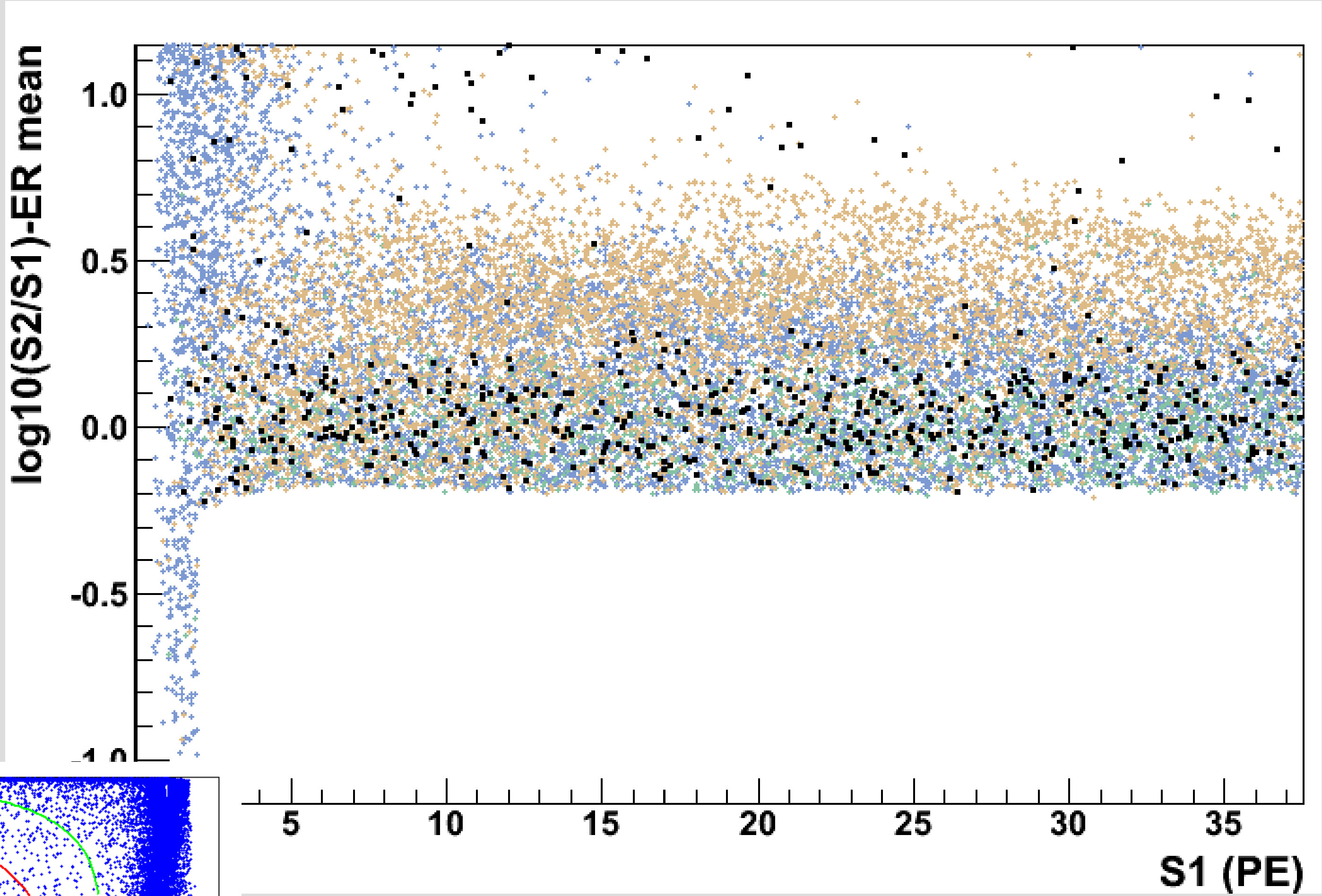


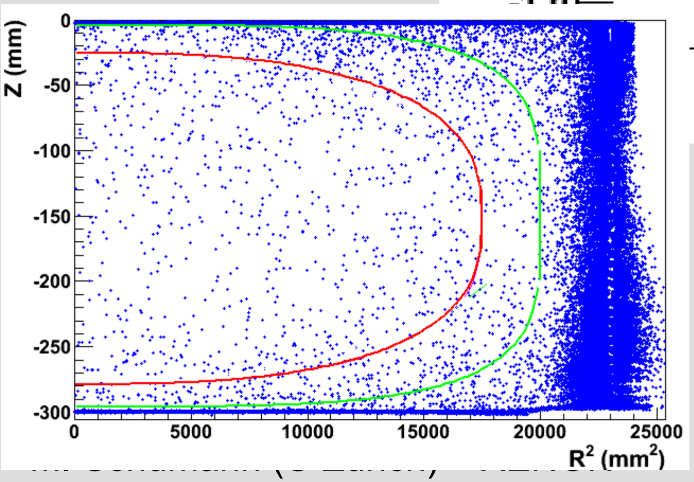
Single Scatter Selection





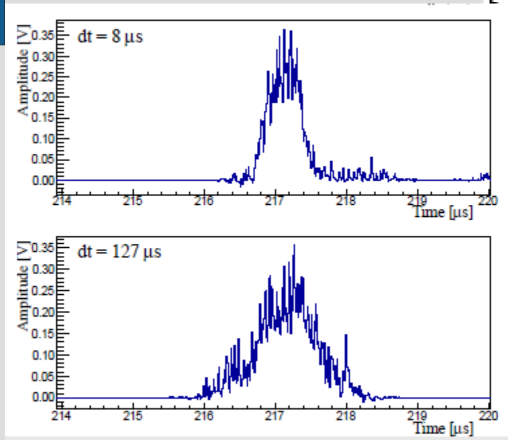
Threshold and Fiducial Volume

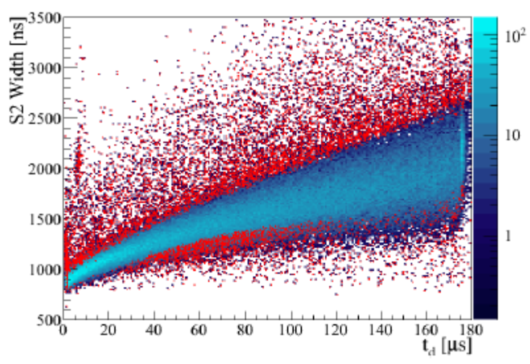


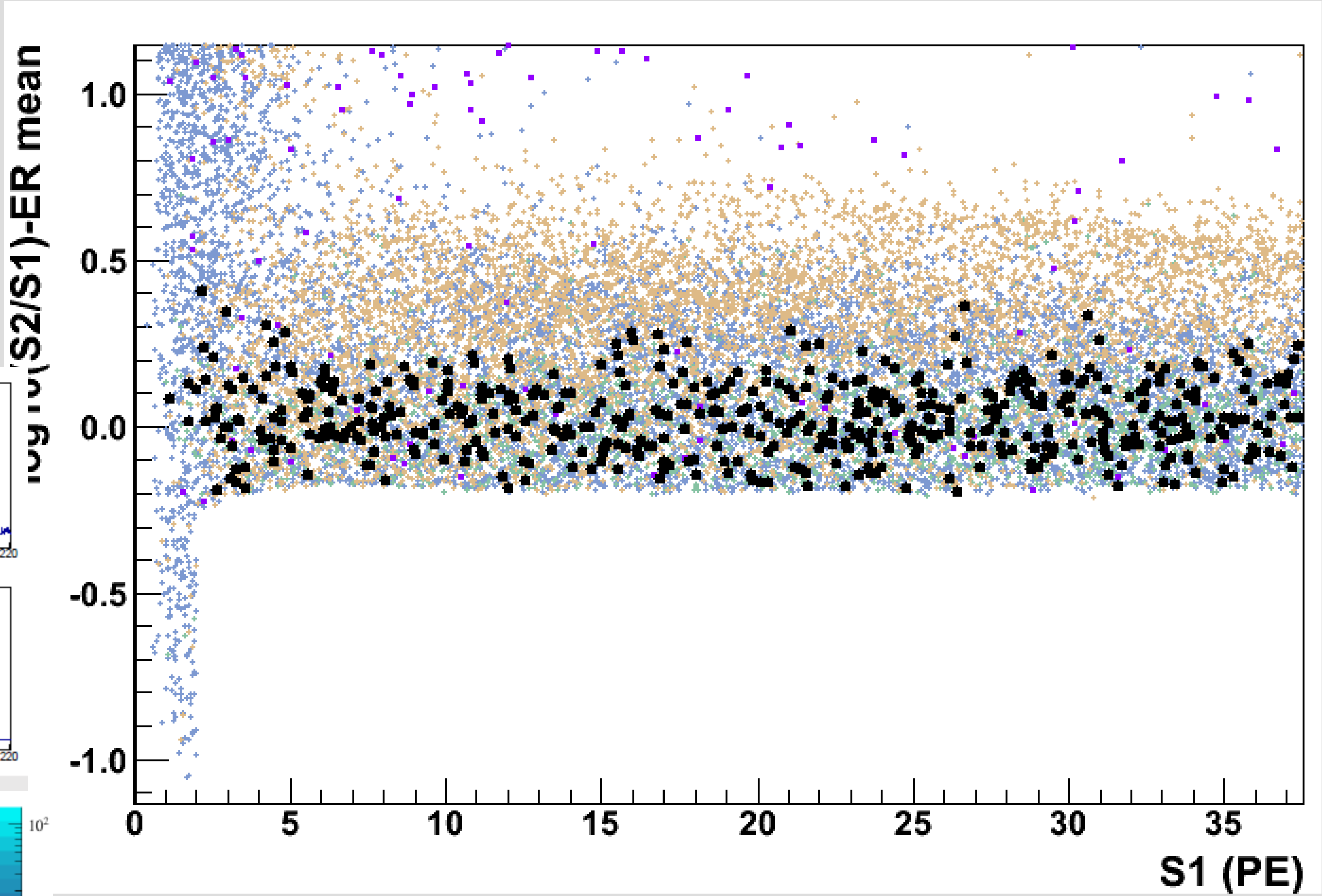


Consistency Cuts

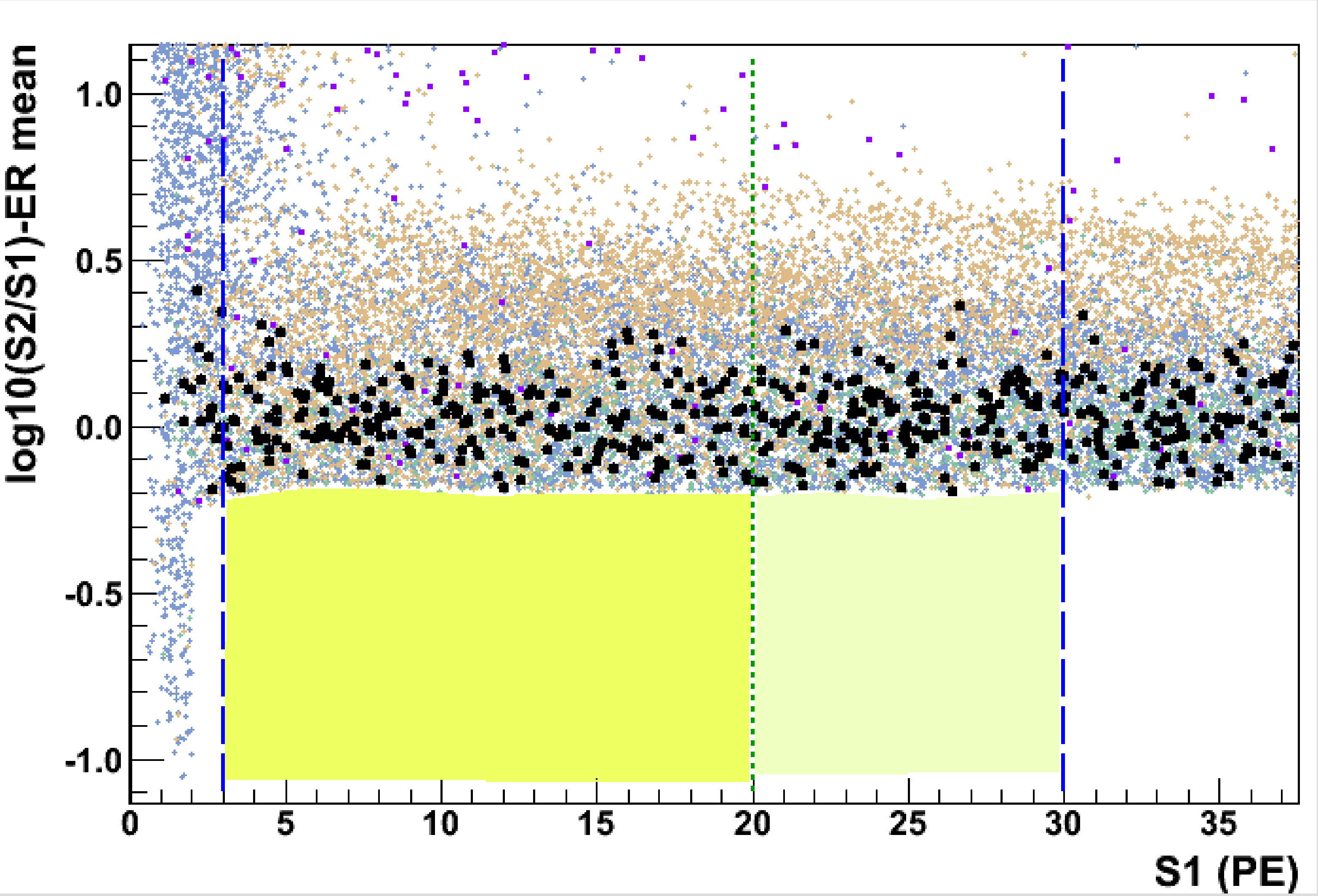
S2 Width:



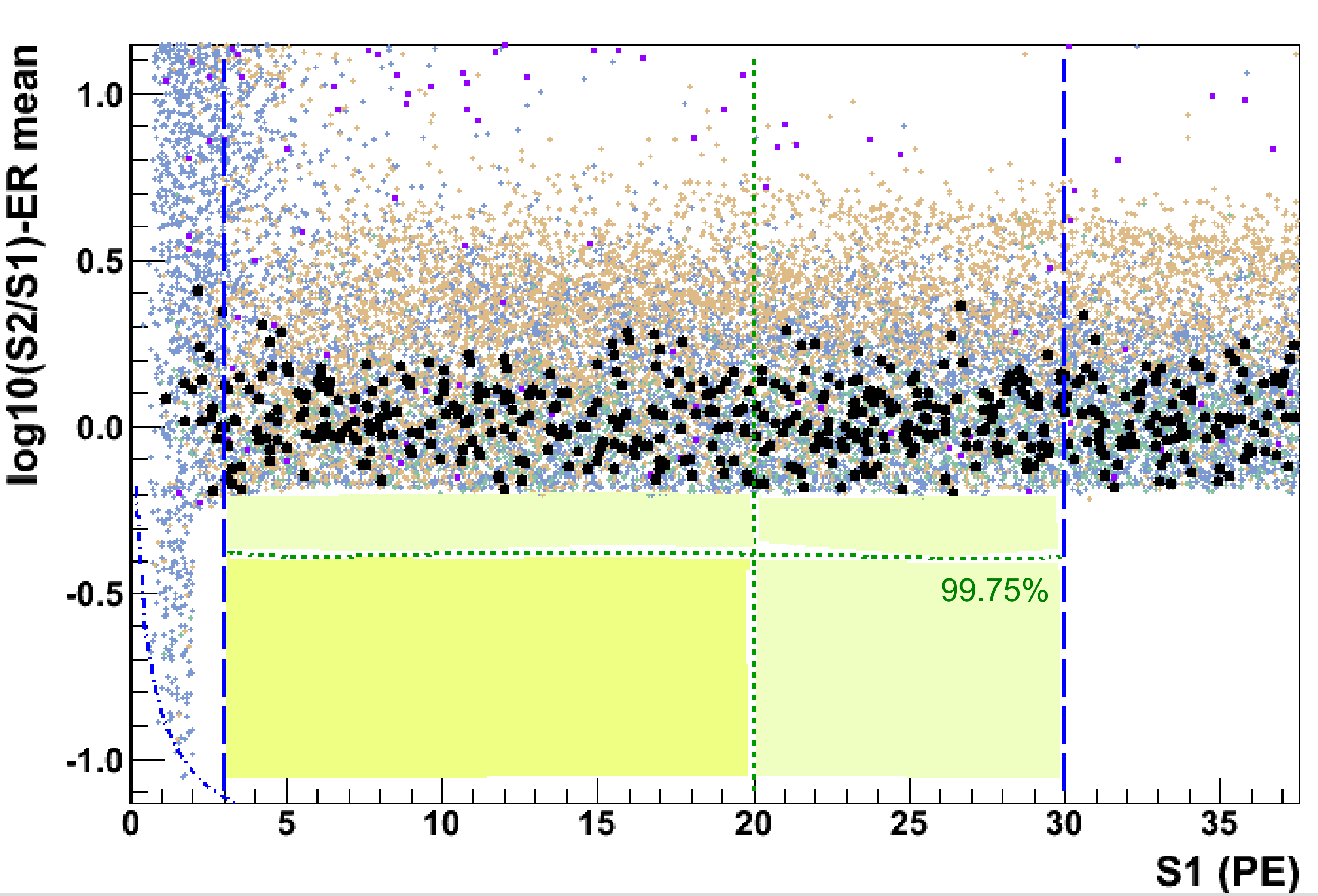




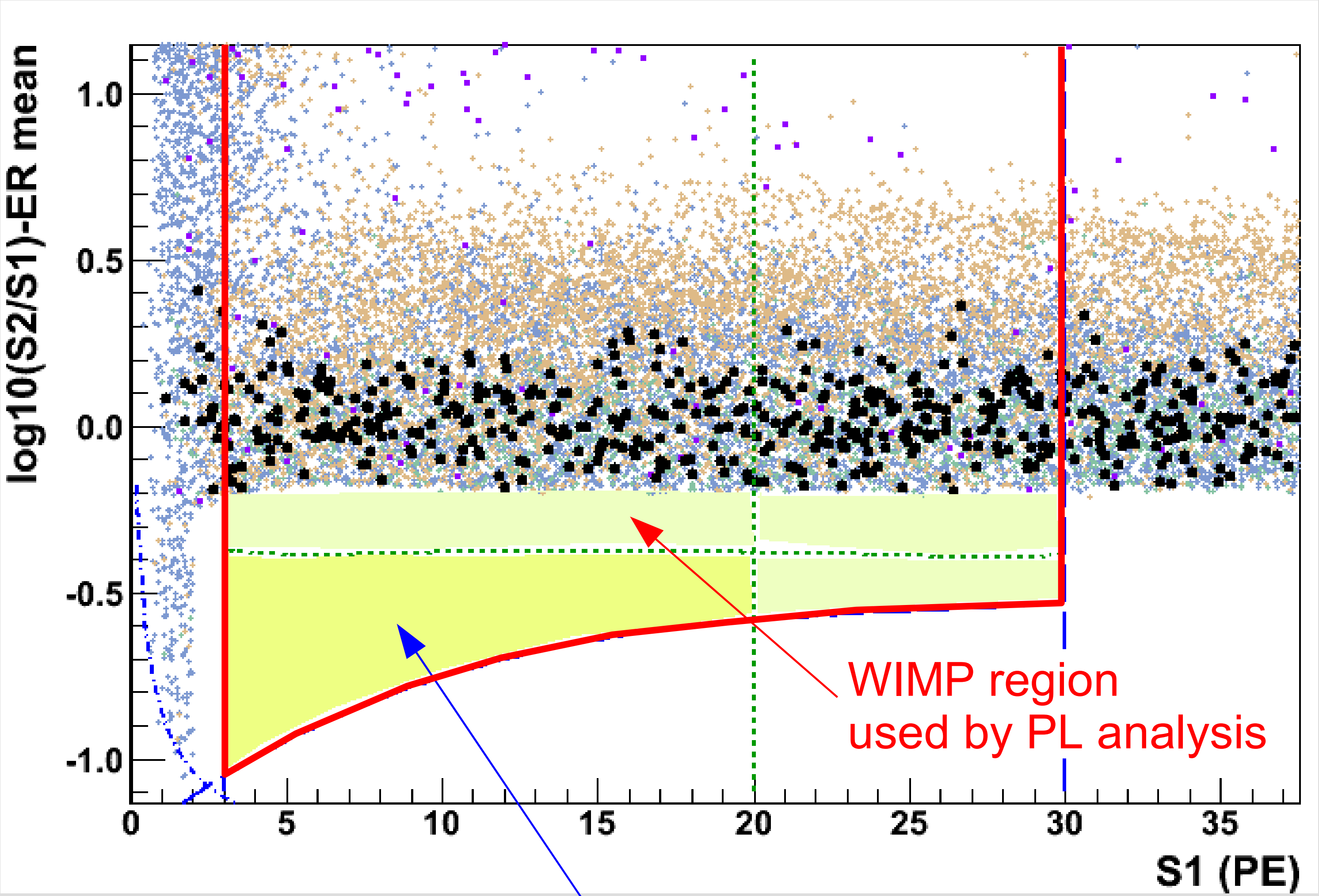
Select Energy Range



ER Rejection



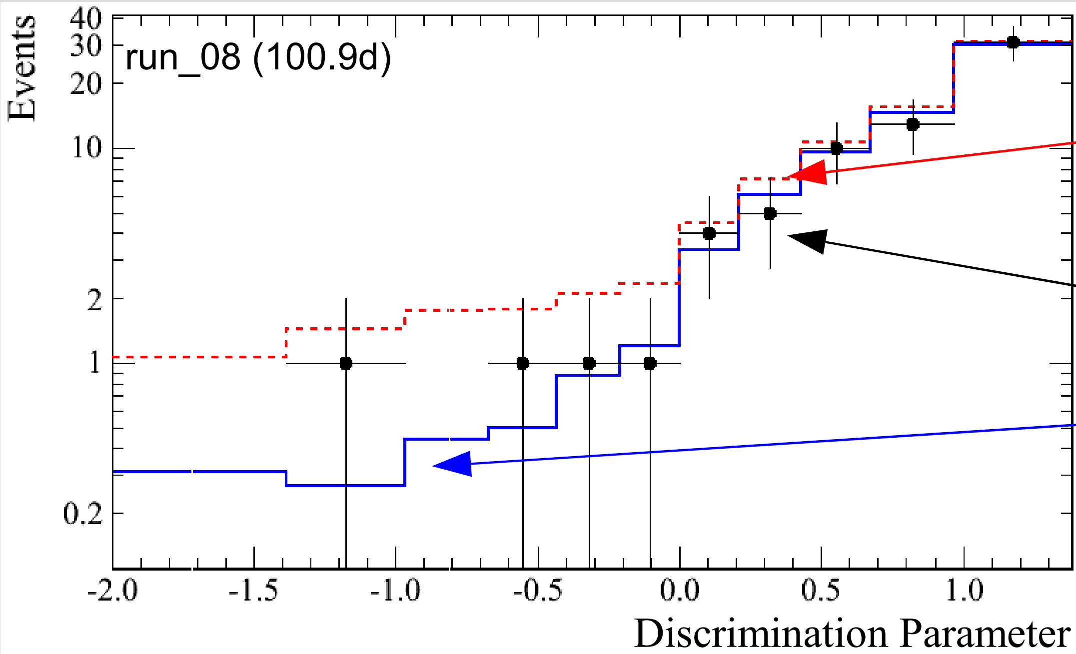
WIMPs are Nuclear Recoil-like



„Benchmark Region“

Profile Likelihood Method

PRD 84, 052003 (2011)



background + WIMP signal
(100 GeV/c² at 10^{-44} cm², 13 events)

observed Signal

expected background

Construct Likelihood function:

$$\mathscr{L} = \mathscr{L}_1(\sigma, N_b, \epsilon_s, \epsilon_b, \mathcal{L}_{\mathrm{eff}}, v_{\mathrm{esc}}; m_\chi) \times \mathscr{L}_2(\epsilon_s) \times \mathscr{L}_3(\epsilon_b) \times \mathscr{L}_4(\mathcal{L}_{\mathrm{eff}})$$

„NR like“ sideband measurement

NR scale measurement

„ER like“ sideband measurement

dark matter measurement

need good understanding of background („background model“)

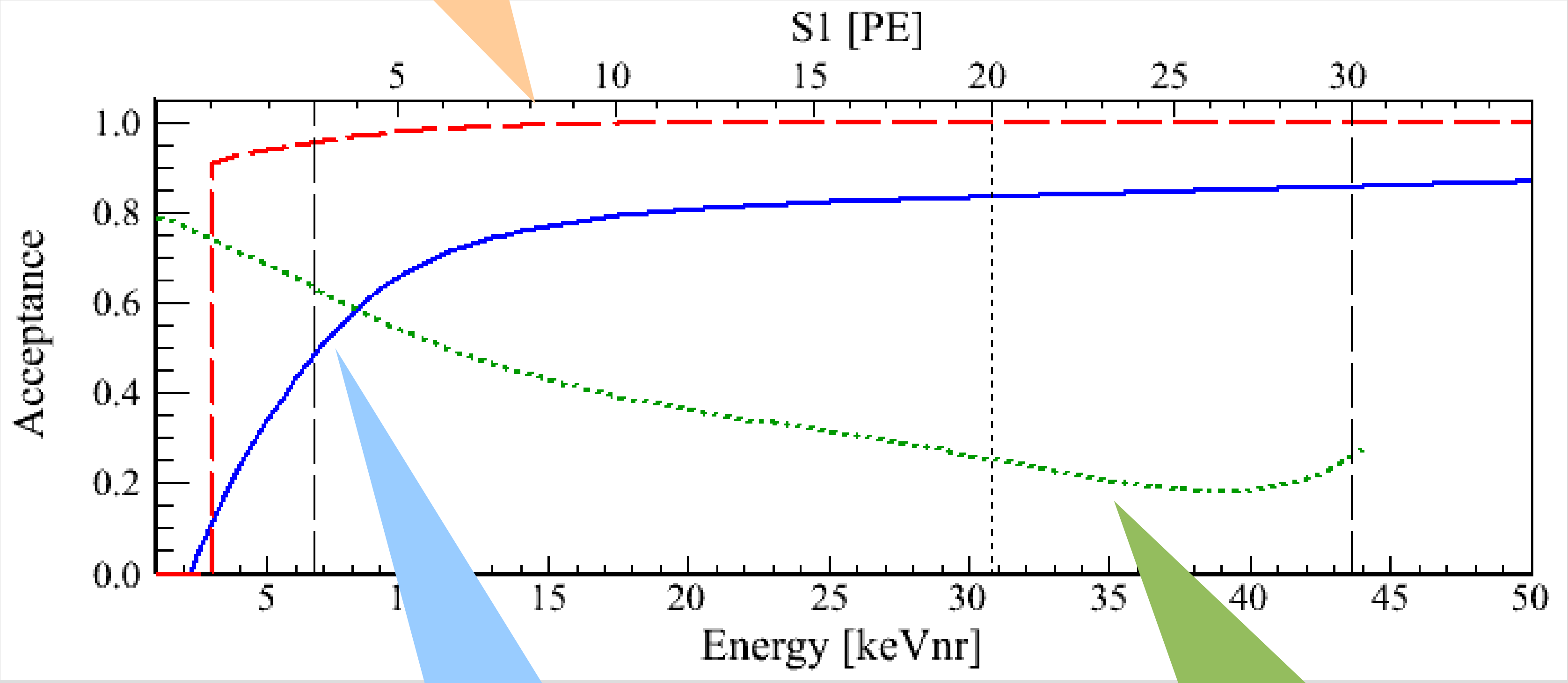
➡ but this is required by any low background experiment (regardless of the type of analysis)

Cuts and Acceptance

S2 Threshold acceptance
(to be applied before S1 smearing)

Erec → Charge → S2
Erec → Light → S1

independent fluctuations



details: *arXiv:1207.3458*

Cuts Acceptance

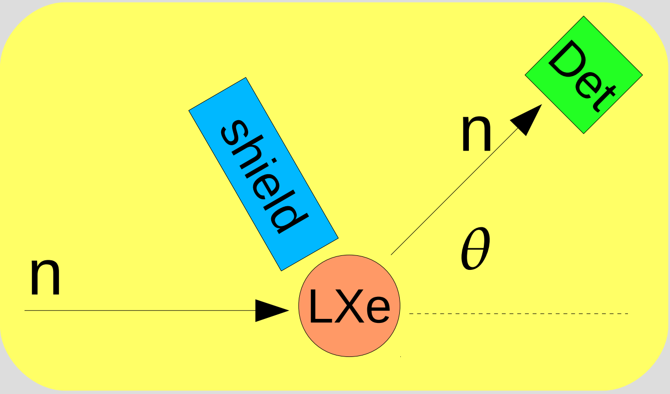
NR acceptance
(benchmark region only)

Nuclear Recoil Energy Scale

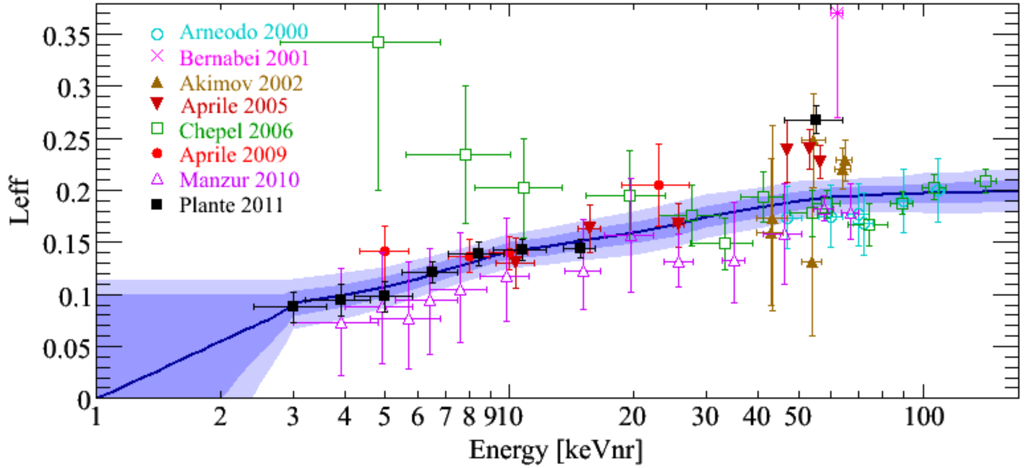
- WIMPs interact with Xe nucleus
  ➞ nuclear recoil (*nr*) scintillation (β and γ's produce electronic recoils)
- absolute measurement of *nr* scintillation yield is difficult
  ➞ measure relative to ^{57}Co (122keV)
- relative scintillation efficiency *L*eff:

$$\mathcal{L}_{\rm eff}(E_{\rm nr}) = \frac{\mathrm{LY}(E_{\rm nr})}{\mathrm{LY}(E_{\rm ee} = 122\ \mathrm{keV})}$$

measurement principle:



average over all direct measurements



most recent measurements:
■ *Plante et al., PRC 84, 045805 (2011)*
△ *Manzur et al., PRC 81, 025808 (2010)*

for discussion of possible systematic errors see *A. Manalaysay, arXiv:1007.3746*

Background Prediction

Neutron background:
- (α,n)+sf and muon induced neutrons
- MC simulation using the exact XENON100 geometry and measured contaminations

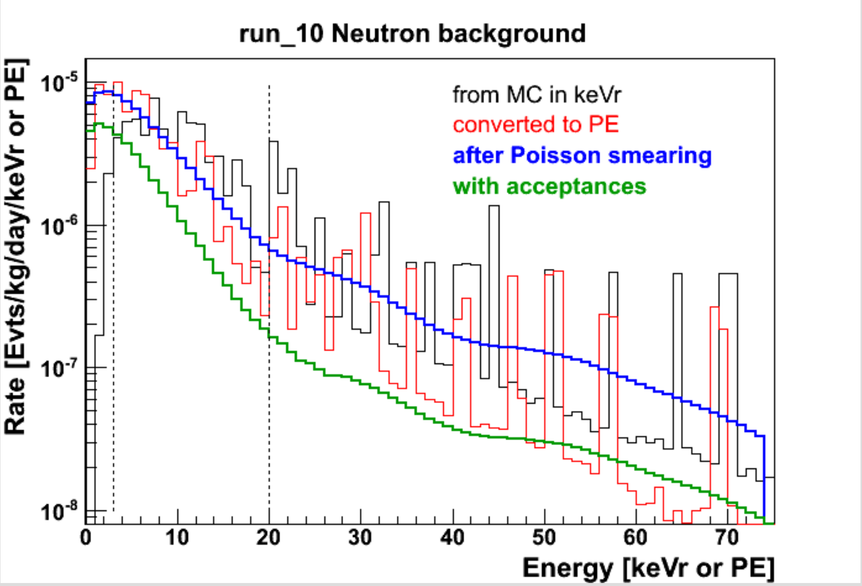
Expect: (0.17 +0.12 –0.07) events

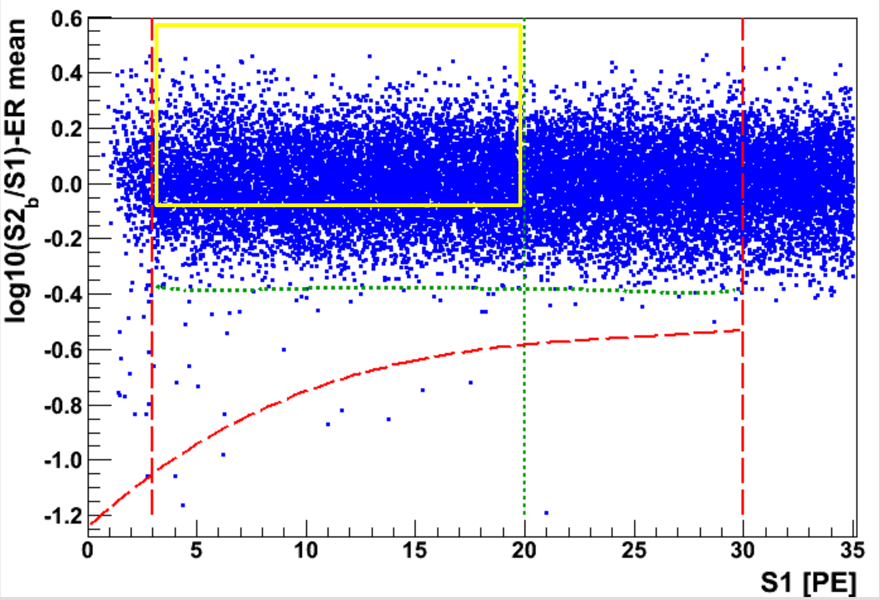
ER background:
- γ activity of the detector and shield
- intrinsic radioactivity in the LXe
  (→ considerably lowered this run)
- use ER calibration to model background by scaling it to the observable DM data

Expect: (0.79 ± 0.16) events

Sum: (1.0 ± 0.2) events

The same background model is implemented in the PL analysis

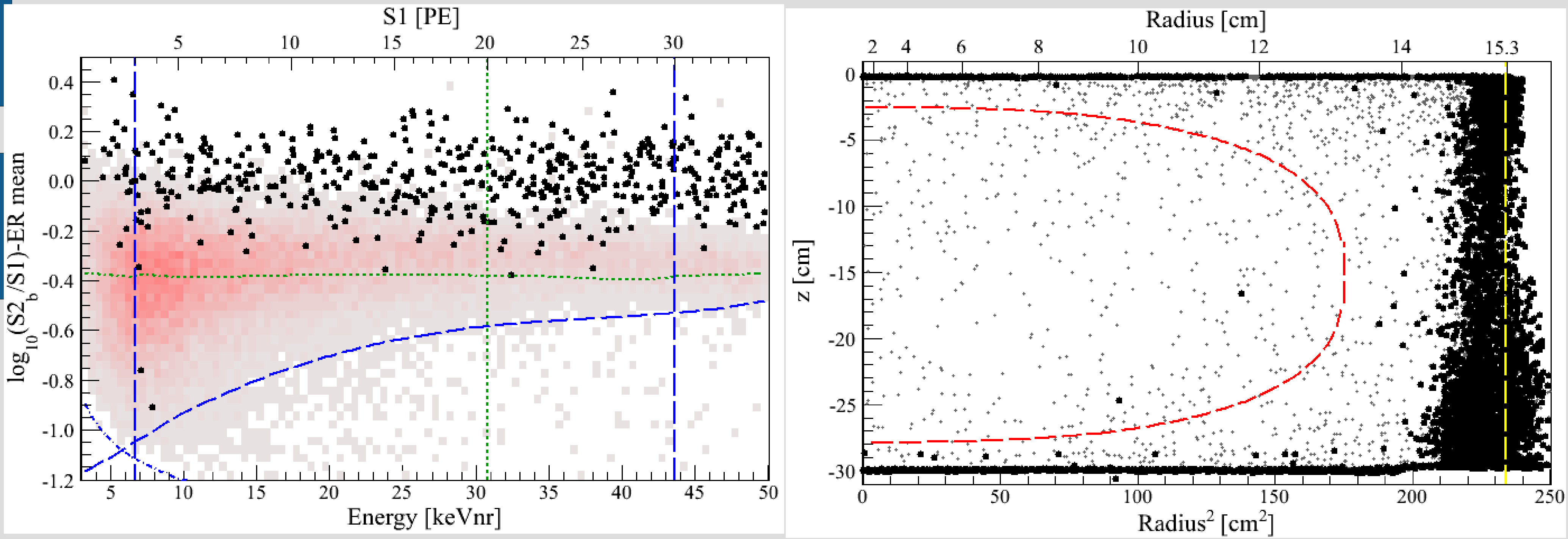




During ...

… and after ...

… Unblinding



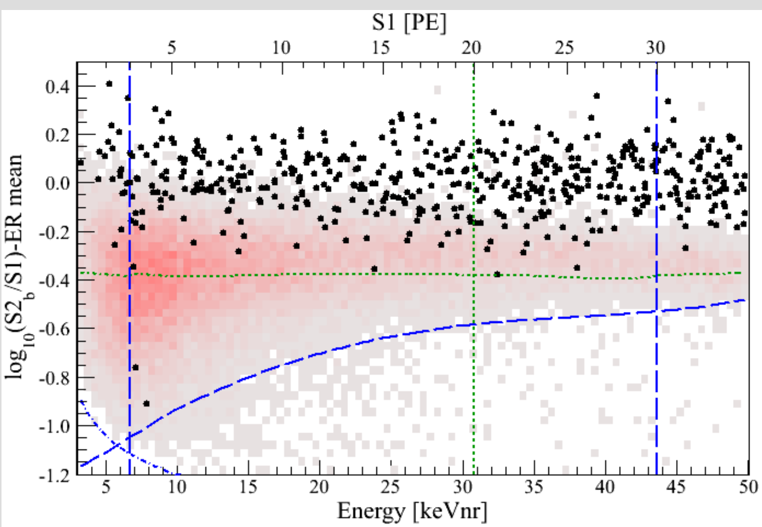
(1.0 ± 0.2) events expected
2 events observed

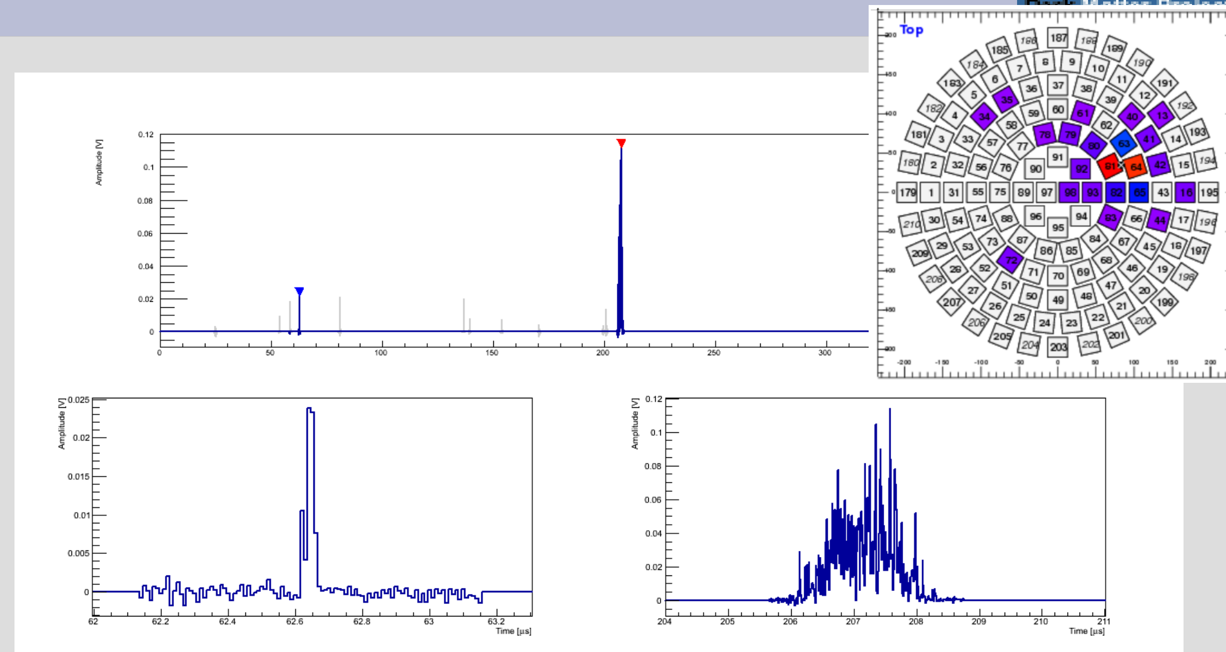
→ 26.4% probability that background fluctuated to 2 events
→ PL analysis cannot reject the background only hypothesis

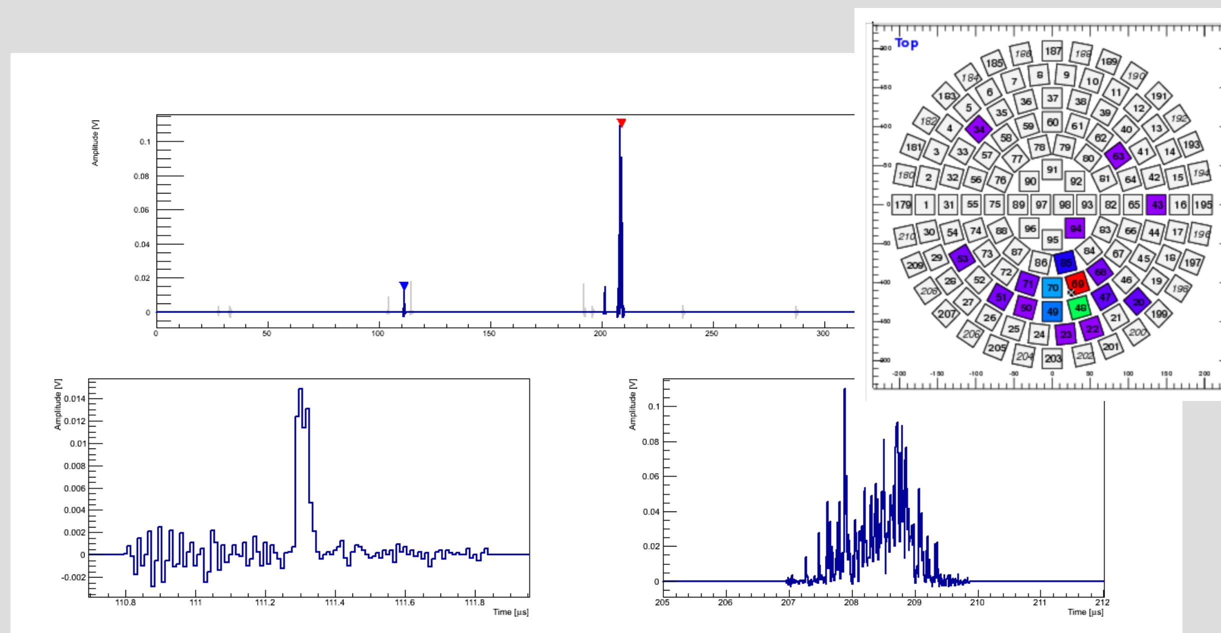
No significant excess due to a signal seen in XENON100 data.

Events in Benchmark Region

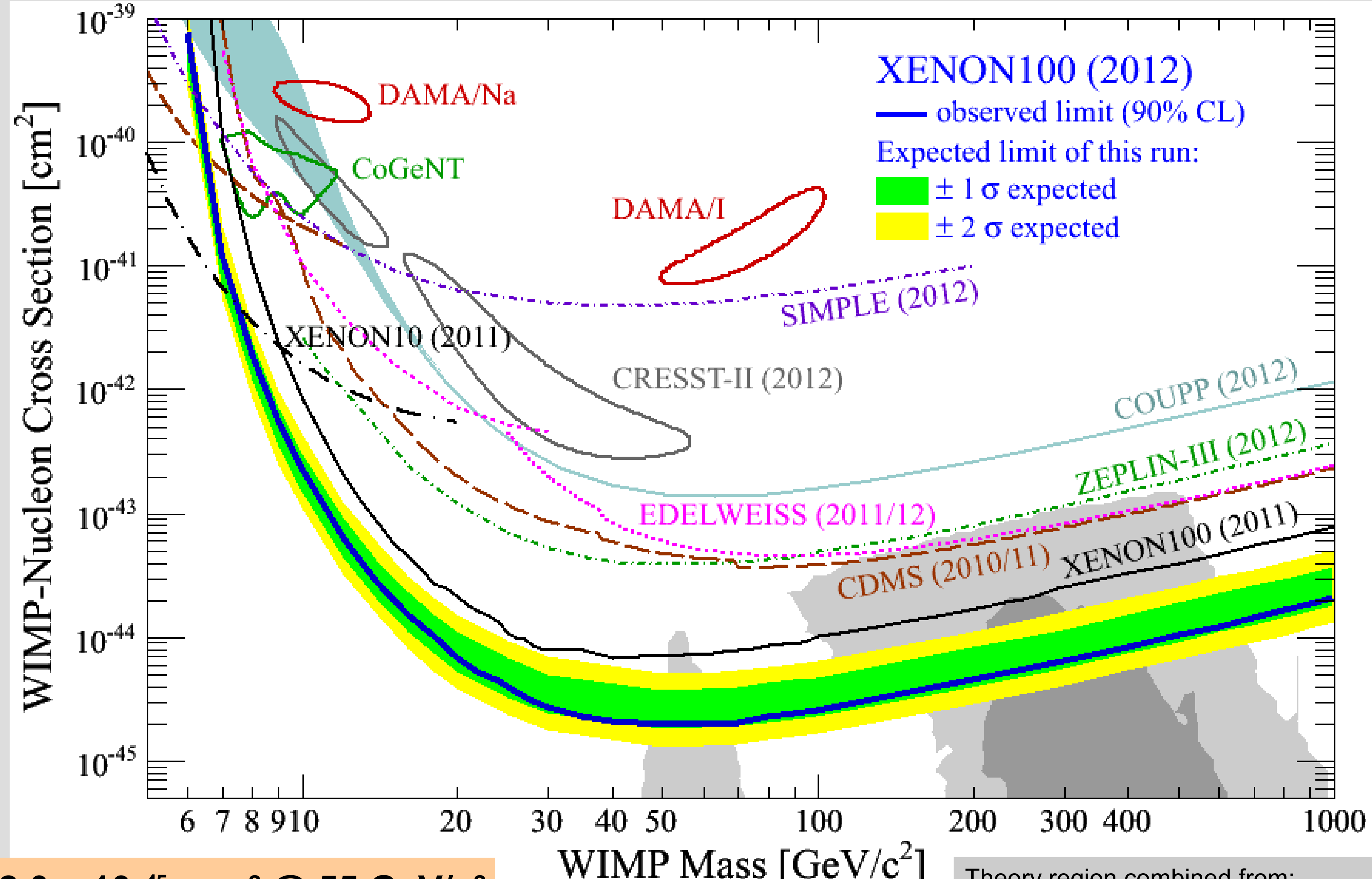
- visual inspection:
  valid waveforms
- at 7.1 keVr and 7.8 keVr
  both events between
  3 and 4 PE
- rather low wrt the
  NR calibration data
- no low S2/S2-events
  below threshold







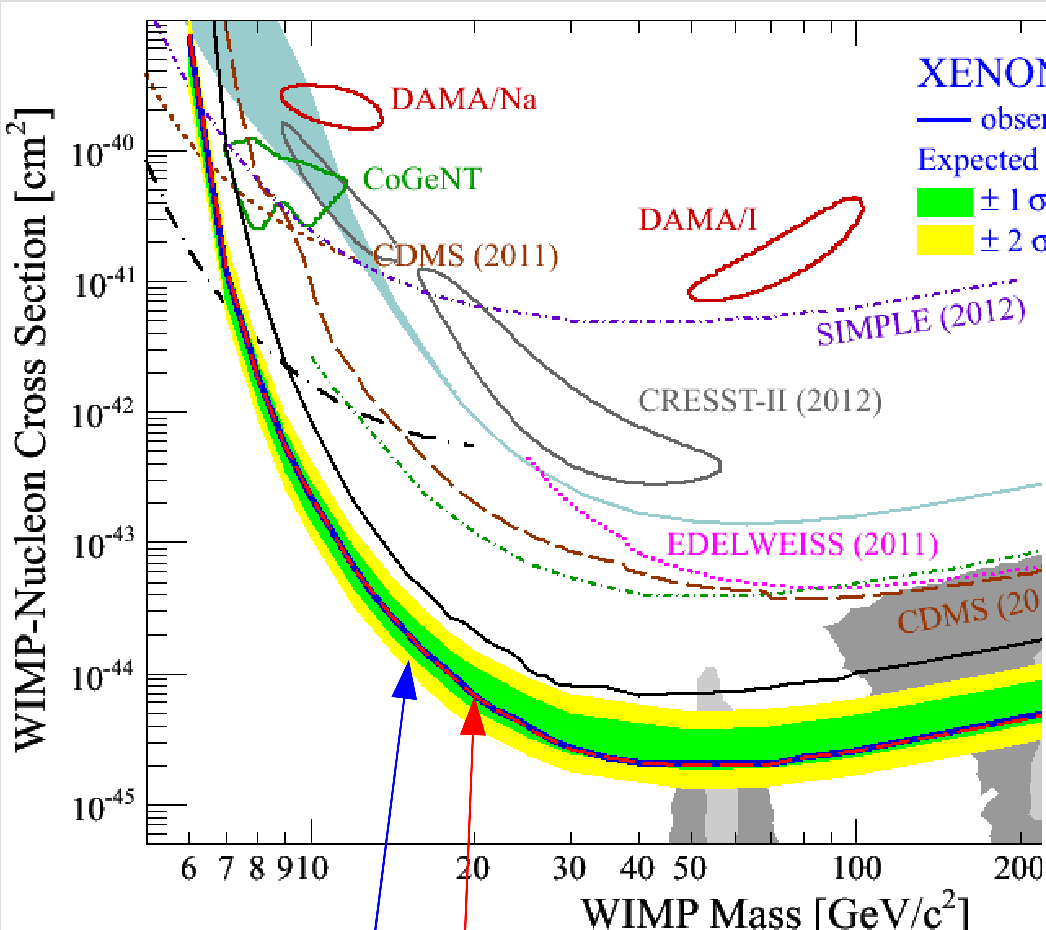
The new XENON100 Limit

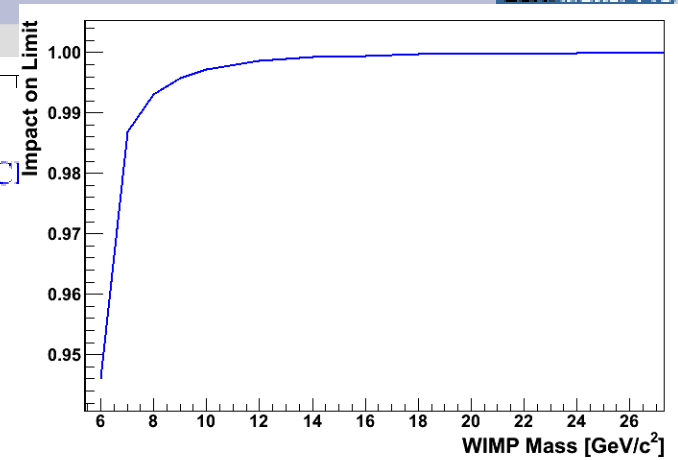


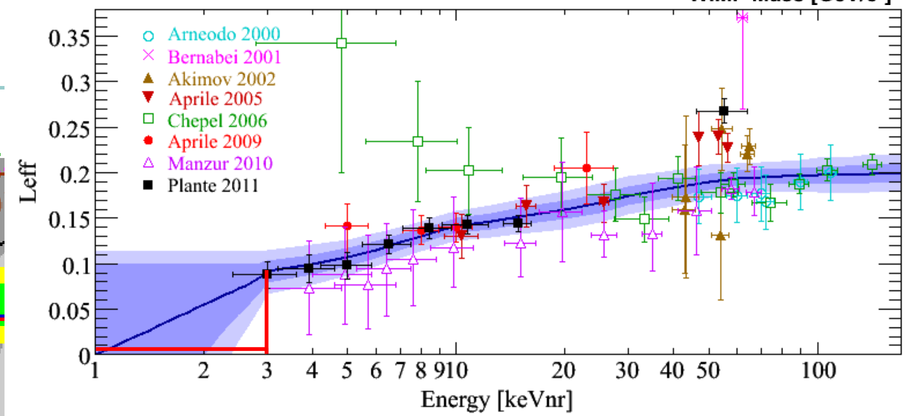
2.0 x 10^{-45} cm² @ 55 GeV/c²

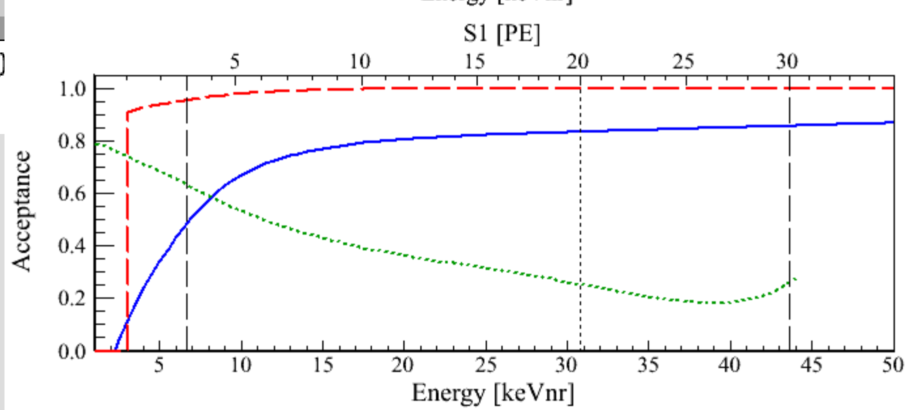
Theory region combined from:
Strege et al., JCAP 1203, 030(2012)
Fowlie et al., arXiv:1206.0264
Buchmueller et al., arXiv:1112.3564

No Impact of L_{eff} below 3 keVr





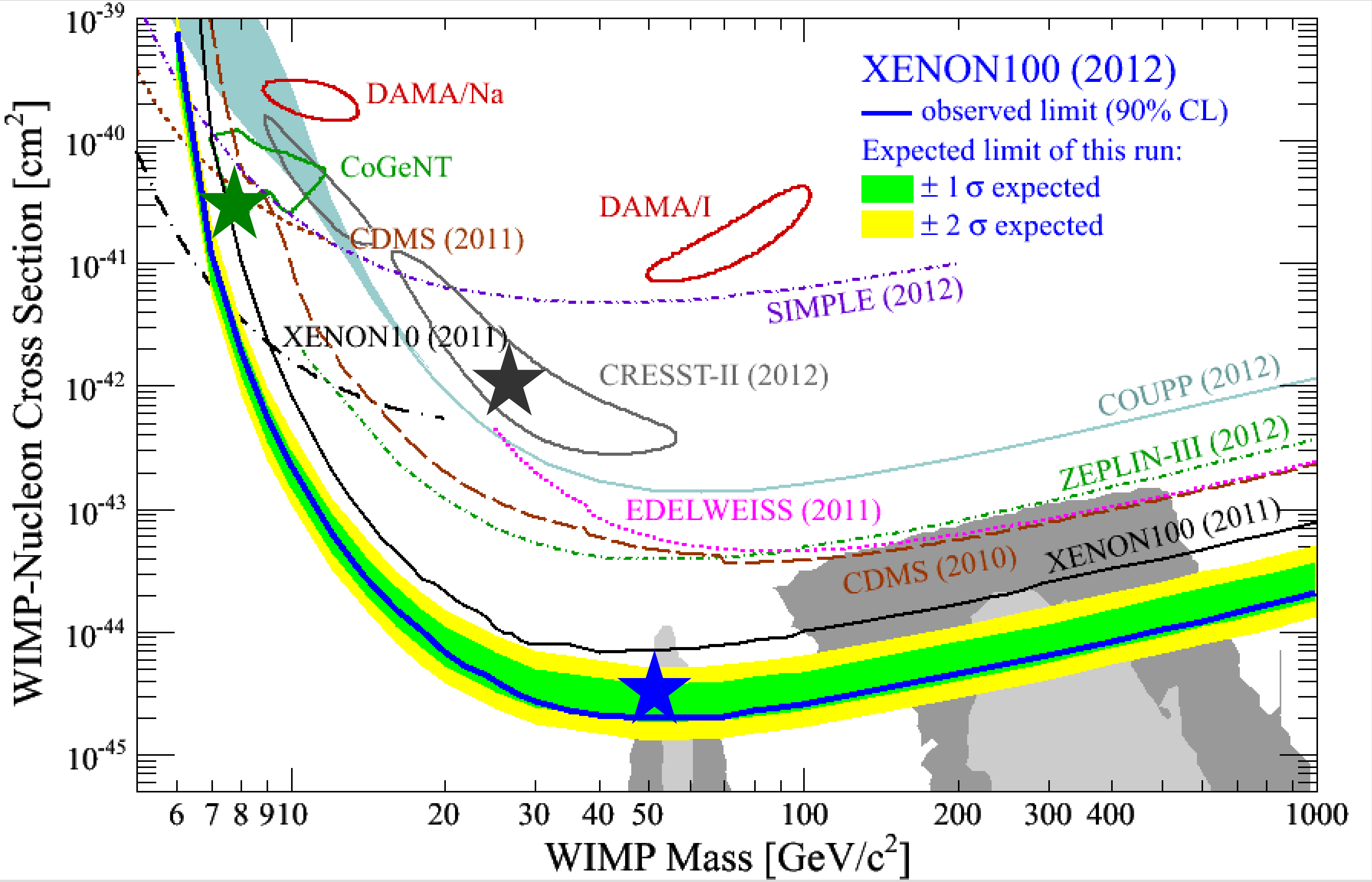




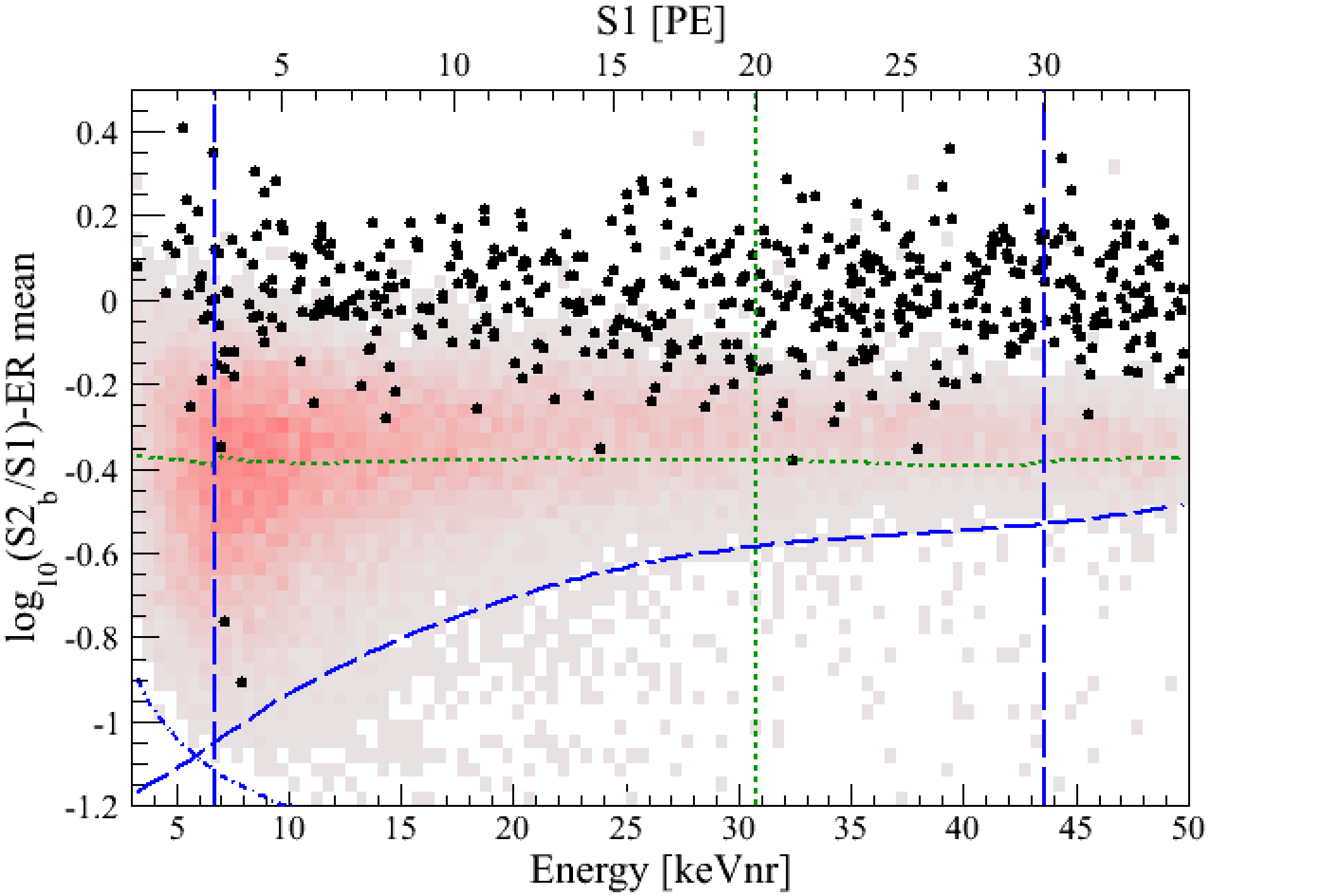
„standard“ Leff

Leff=0 below 3 keVnr

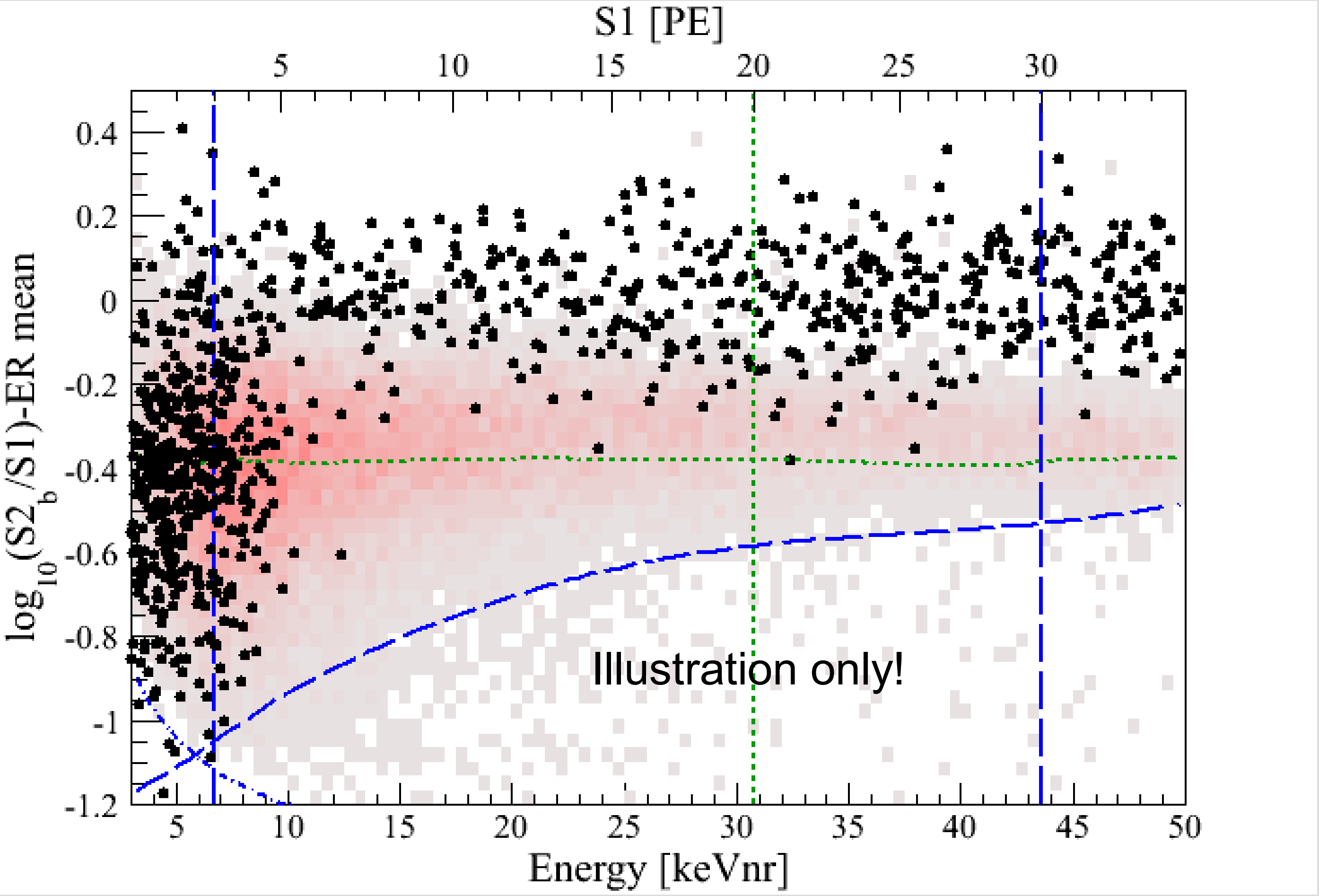
The new XENON100 Limit



What XENON100 sees...

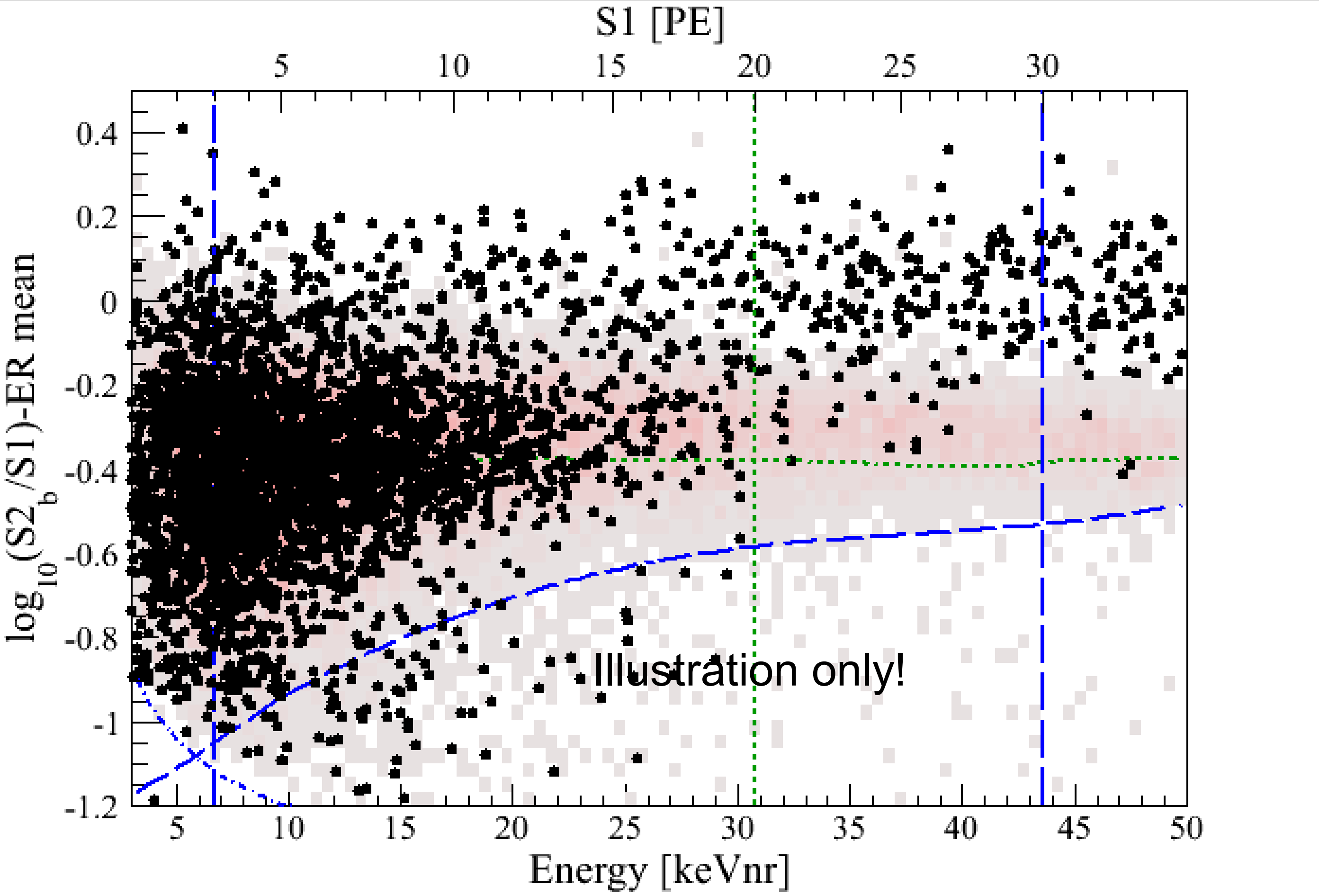


A light mass WIMP...



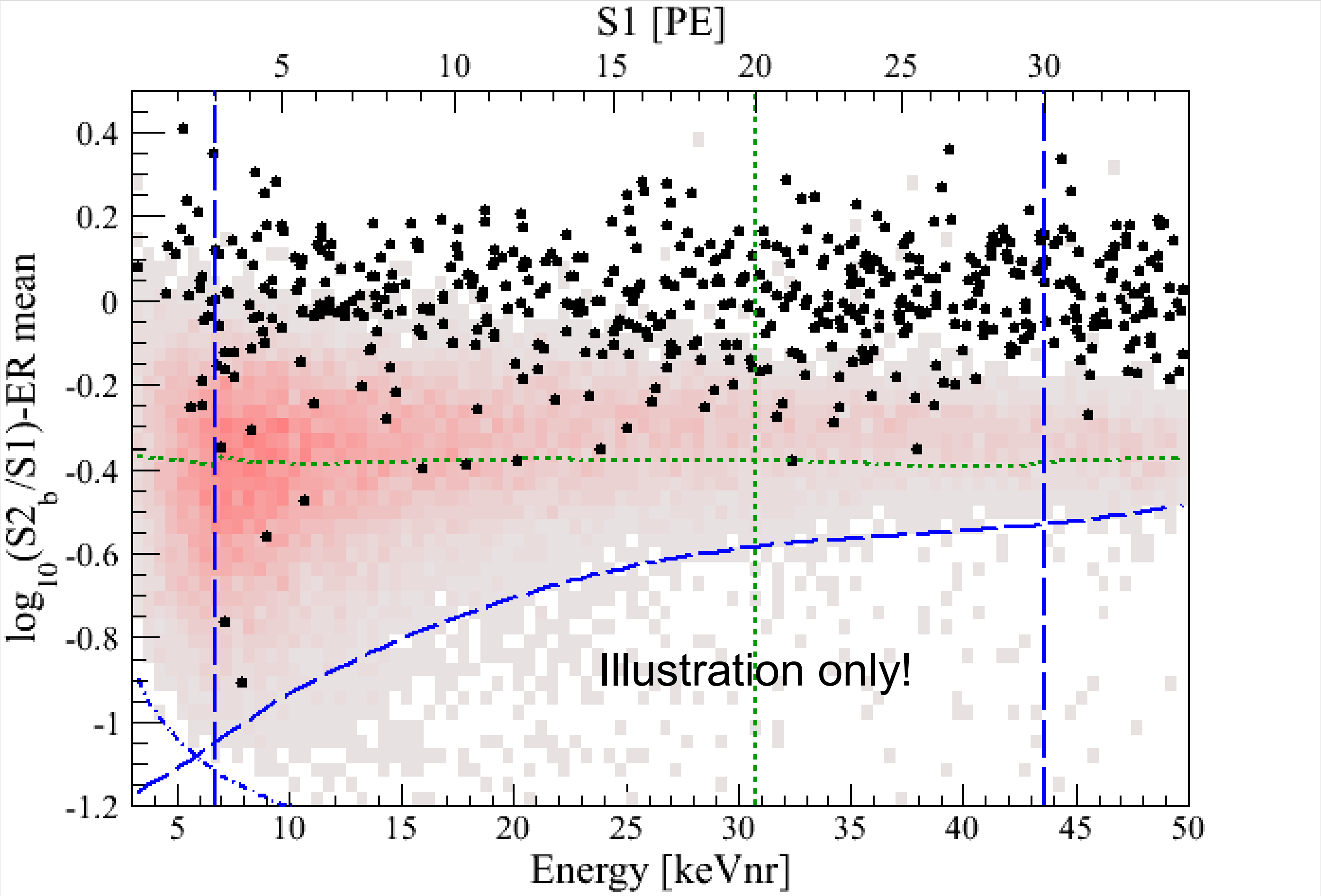
m_χ= 8 GeV/c² σ=3.0x10^{-41} cm²

A CRESST-like signal...



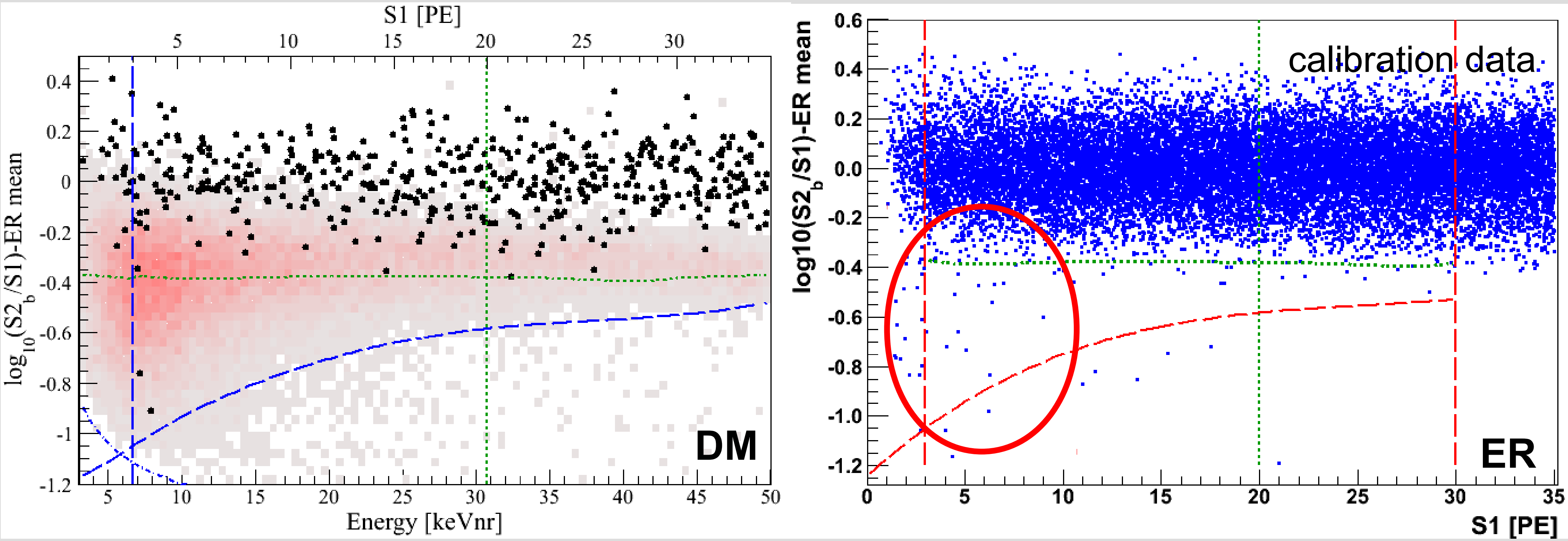
m_χ = 25 GeV/c² σ=1.6 x10^{-42} cm²

What XENON100 excludes...



m_χ= 50 GeV/c² σ=3.0 x10^{-45} cm²

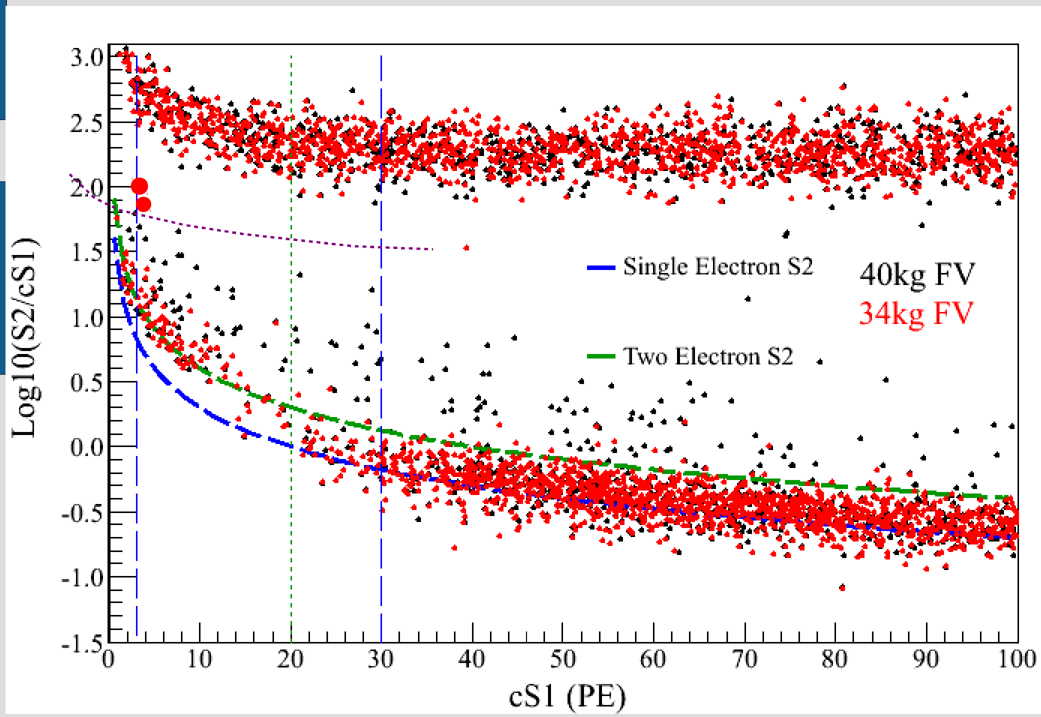
What could the Events be?

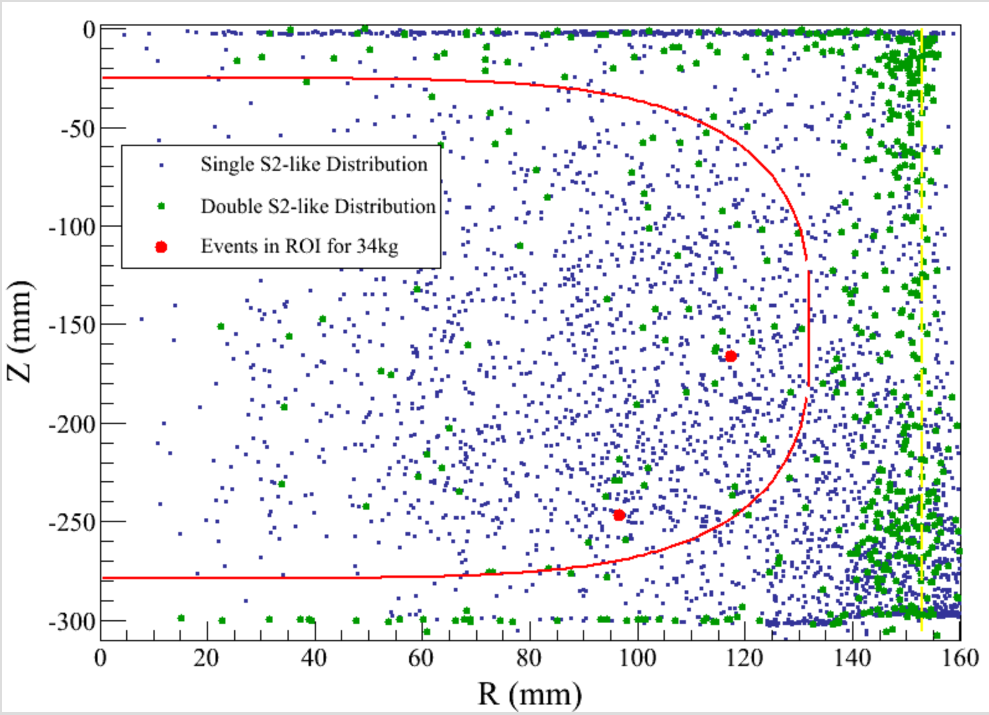


Reminder:
Background is modeled using ER calibration data from Co60 and Th232
This data shows an increased probability for anomalous leakage below ~8 PE

Background prediction depends on the information which is put into the model

Sensitivity to single electrons

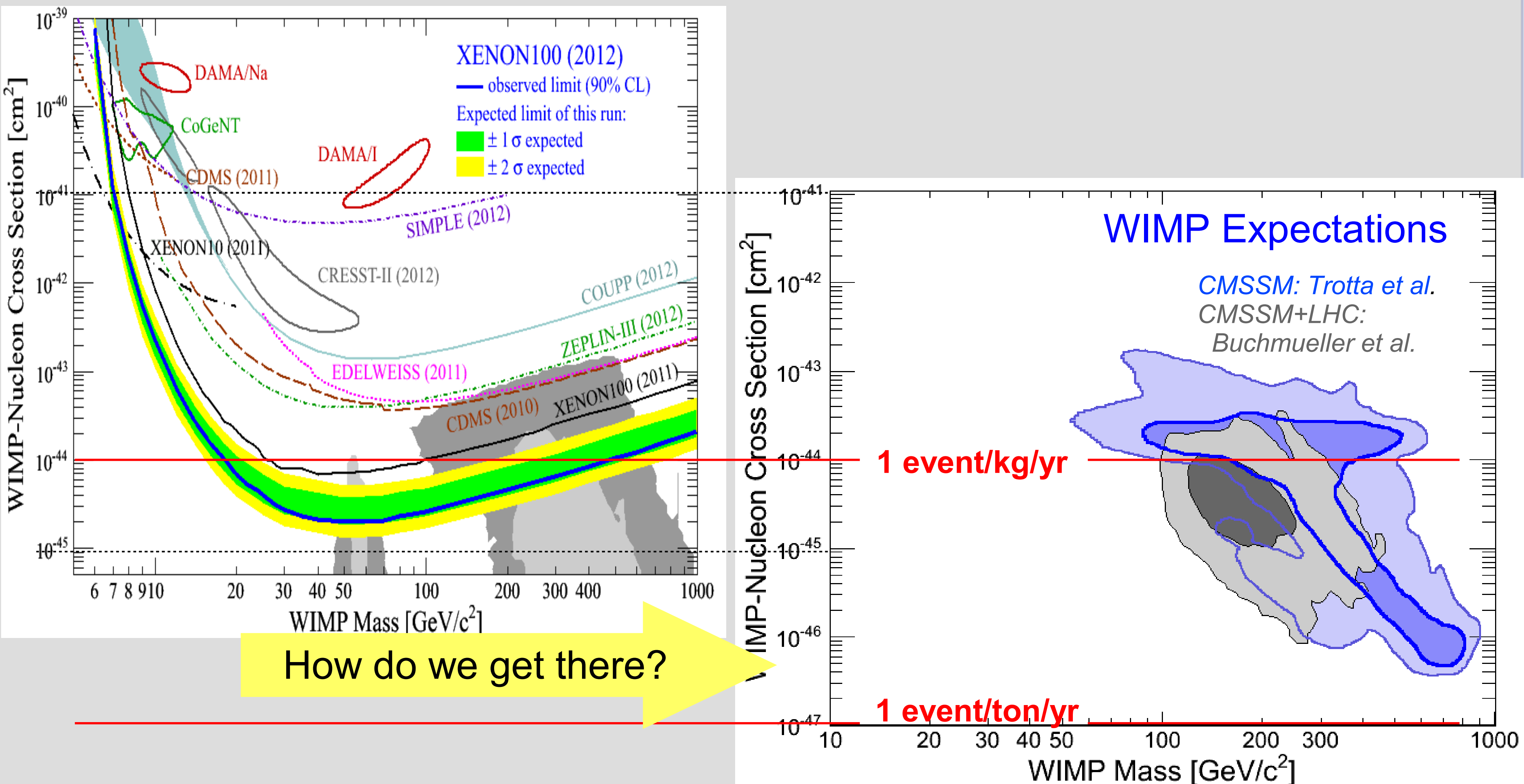




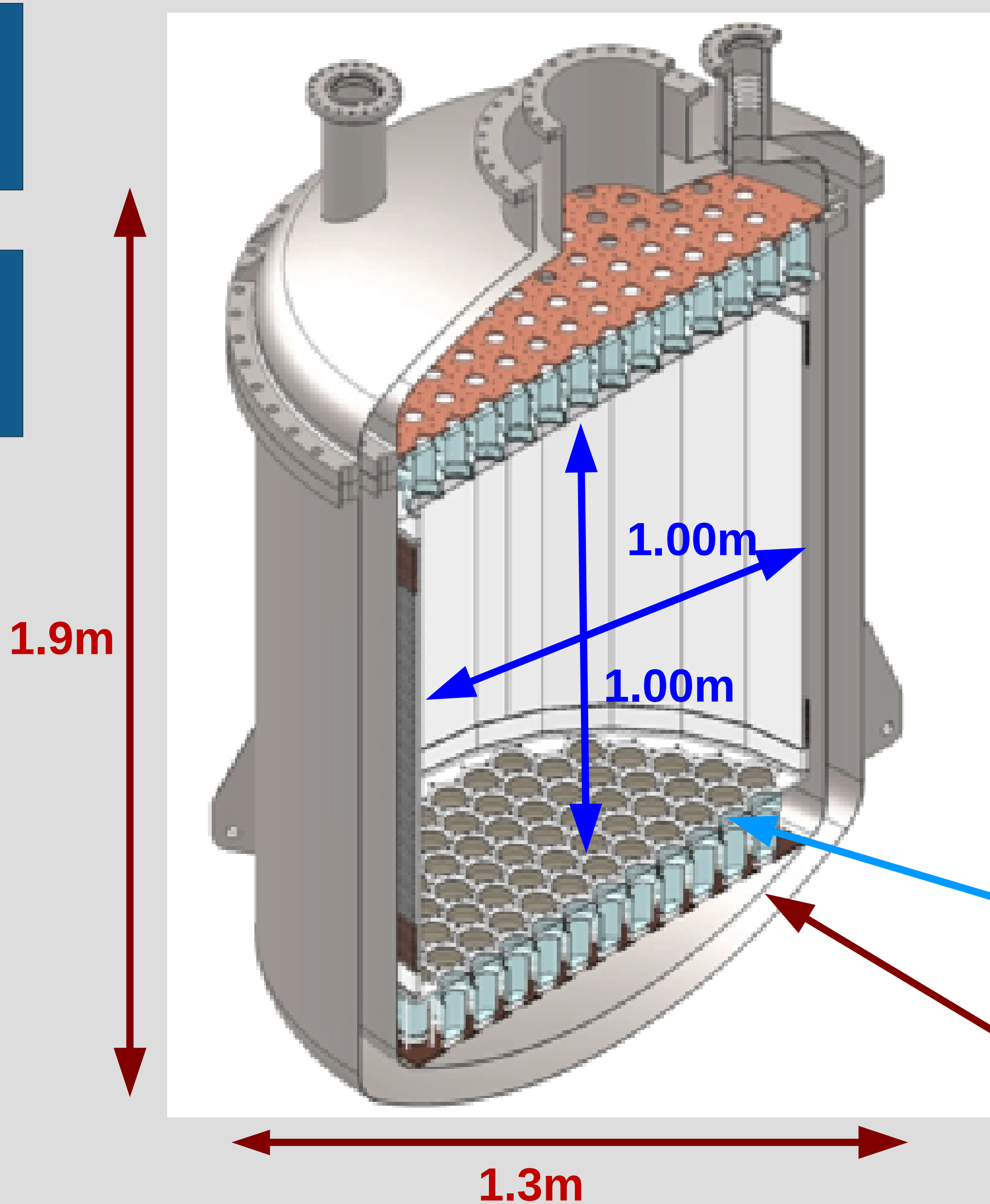
Relaxing the S2 threshold condition (S2>150 PE)
leads to a band of events at very low S2/S1(below signal range)

→ can the 2 events be in the tail of this band???
→ further studies are required
→ aim: quantify and put into background model for the next run

XENON: What next?



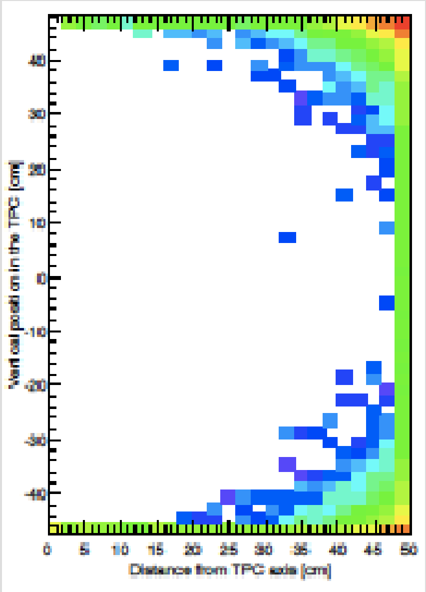
The next step: XENON1T



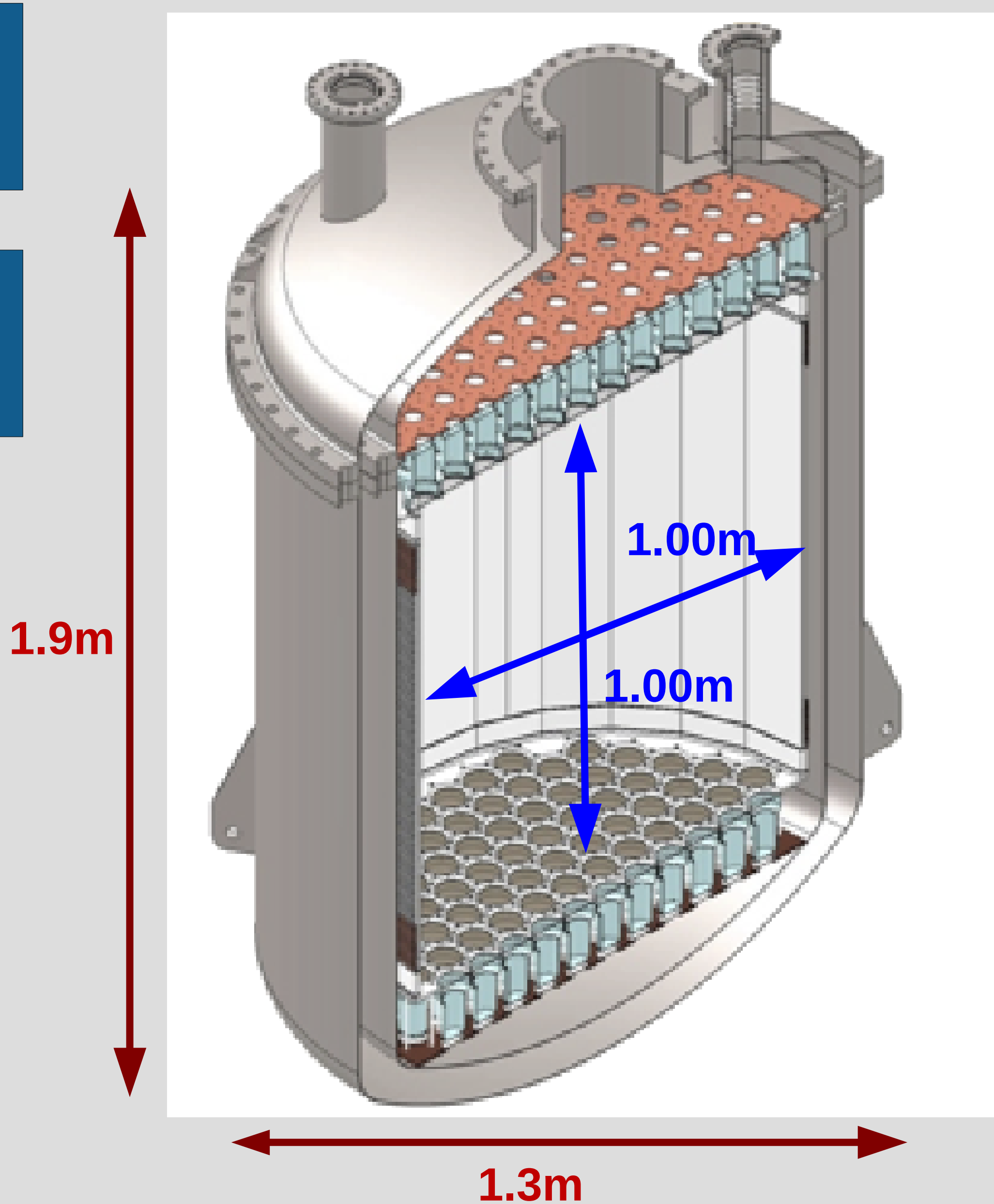
- 3t LXe ("1m³ detector")
  1t fiducial mass → 20x larger
- 100x lower background
  (~10 cm self shielding,
  low radioactivity components)
- background goal: <1 evt in 2 years

Low Radioactivity
Photon Detectors
(3", Total ~250)

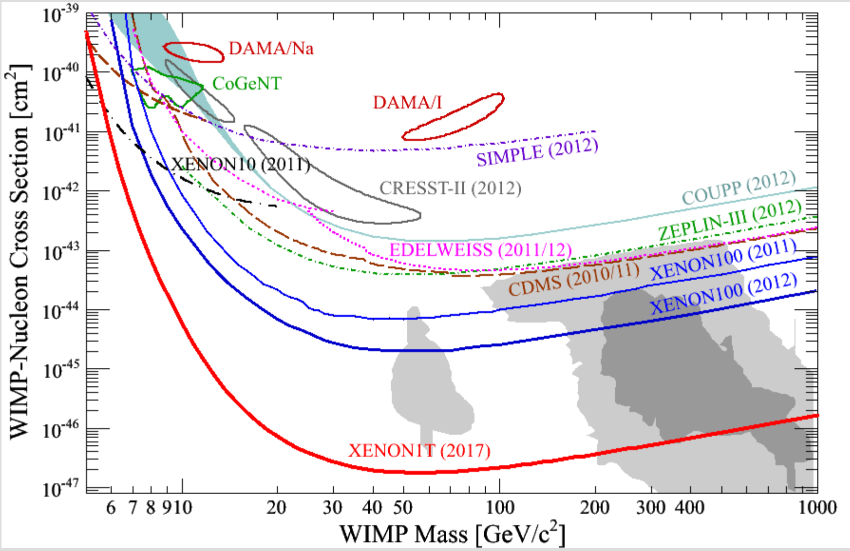
low radioactivity
stainless steel
cryostat
(or copper)



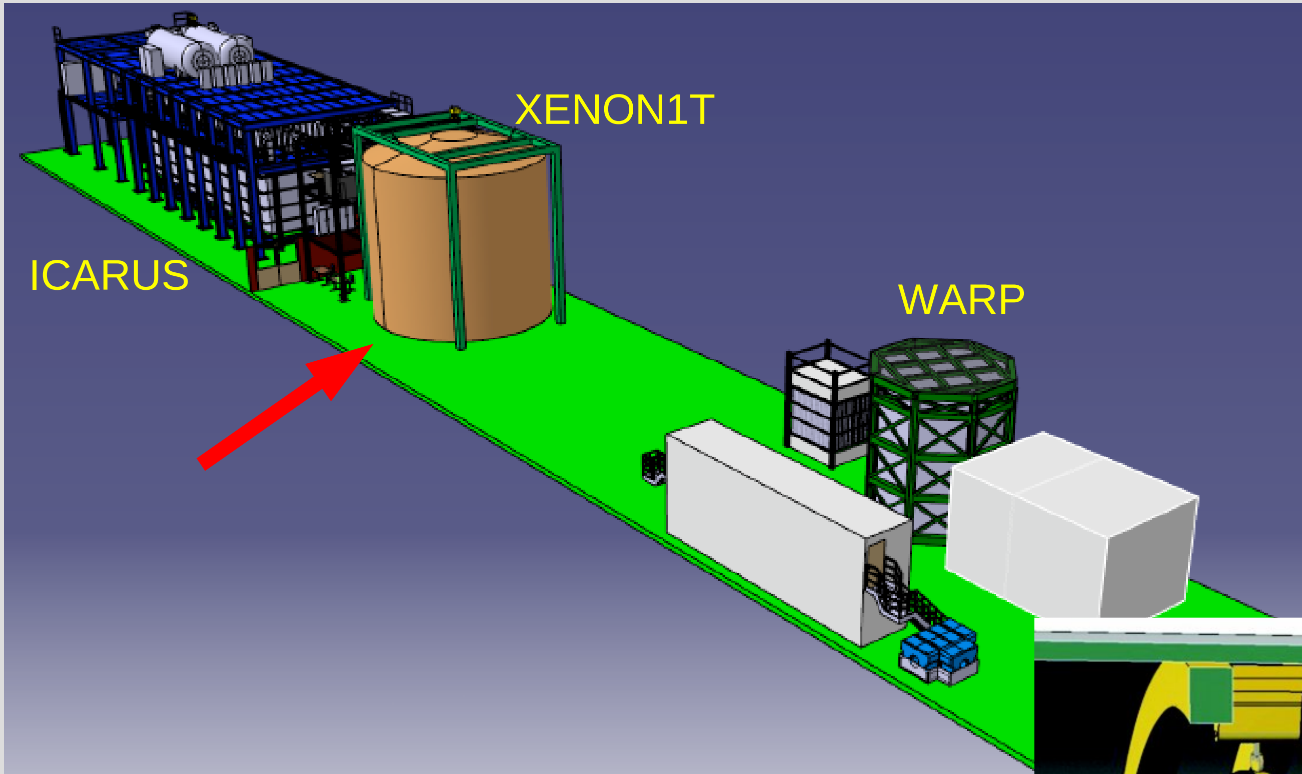
The next step: XENON1T



- 3t LXe ("1m³ detector")
  1t fiducial mass → 20x larger
- 100x lower background
  (~10 cm self shielding,
  low radioactivity components)
- background goal: <1 evt in 2 years
- Timeline: 2010 – 2017
- start construction early 2013

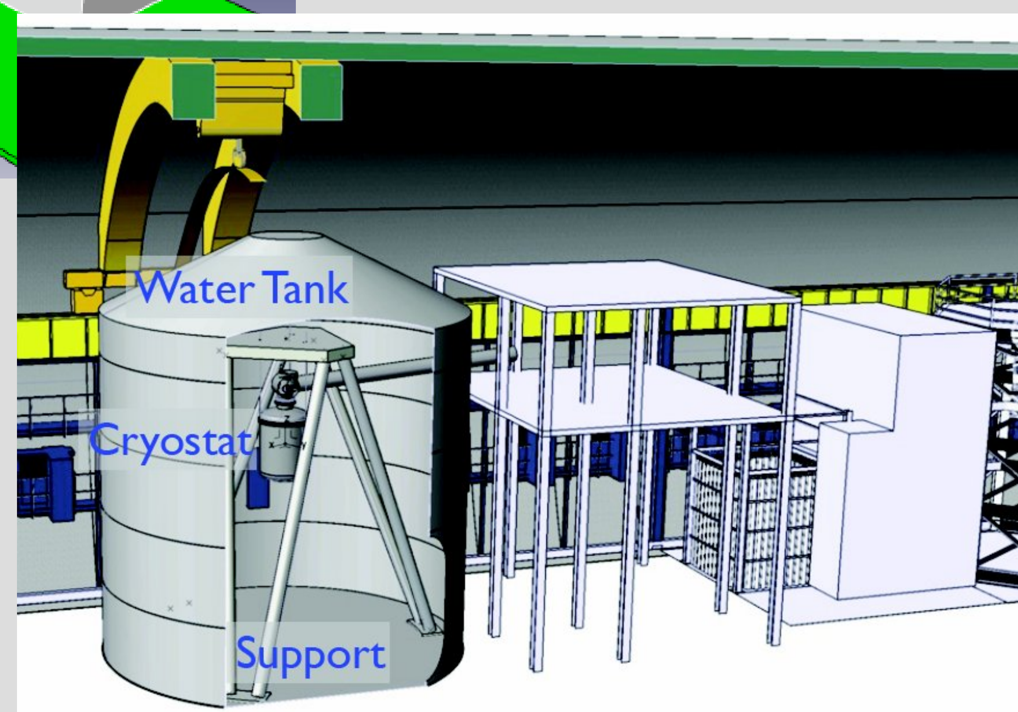


XENON1T @ LNGS



XENON1T
@ LNGS (Hall B)
➞ 4.8 m radius water shield acting as active muon veto

- Proposal and TDR submitted to LNGS
- Approved by INFN end of April 2011
- Approved by NSF (US) May 2012



The new WIMP Landscape